



University of Belgrade
Faculty of Mathematics
Department of Astronomy



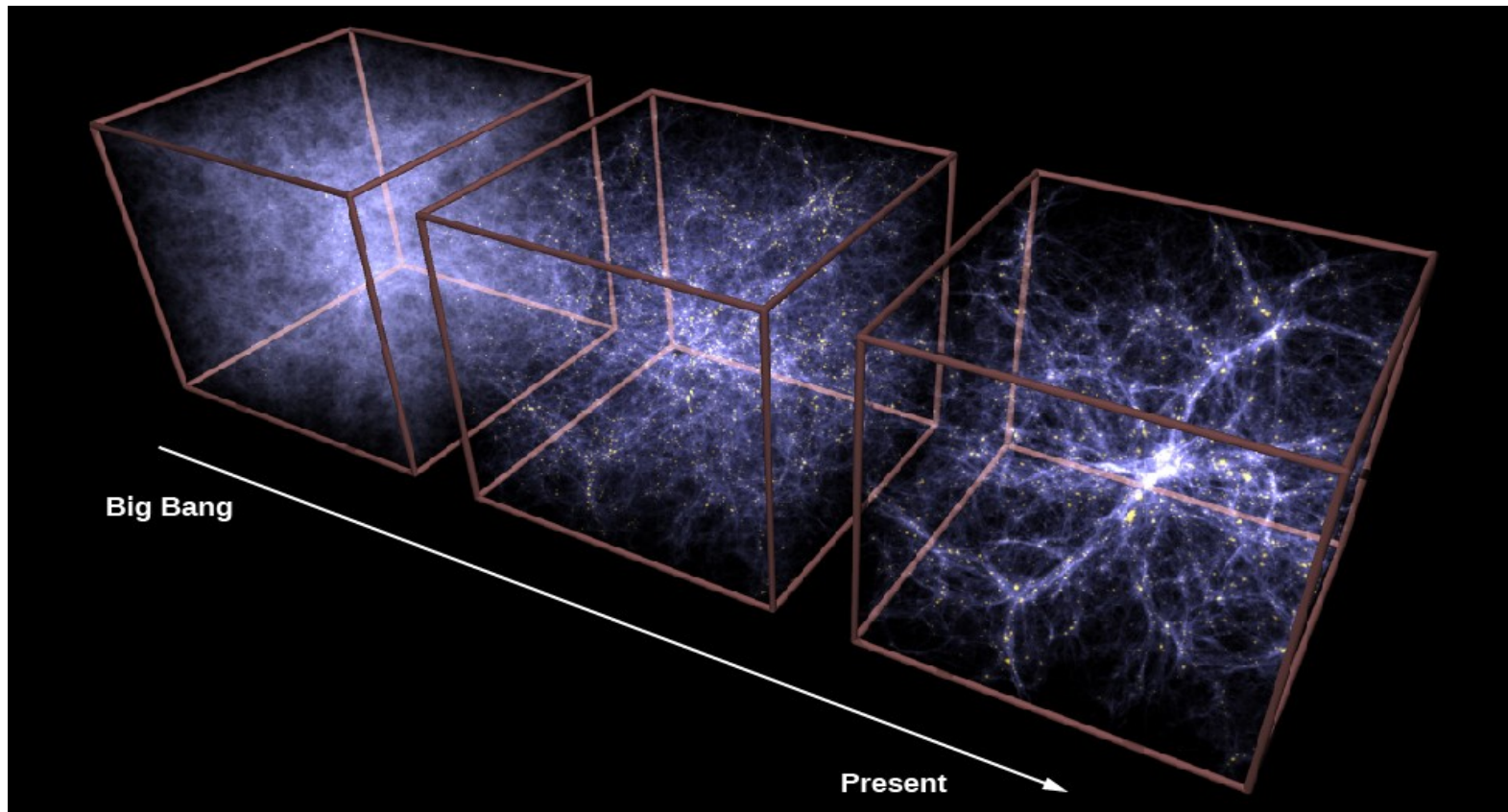
Dark matter halos in galaxy mergers

Stanislav Milošević
stanislav@matf.bg.ac.rs

XIII BSAC
Velingrad, Bulgaria
October, 2022

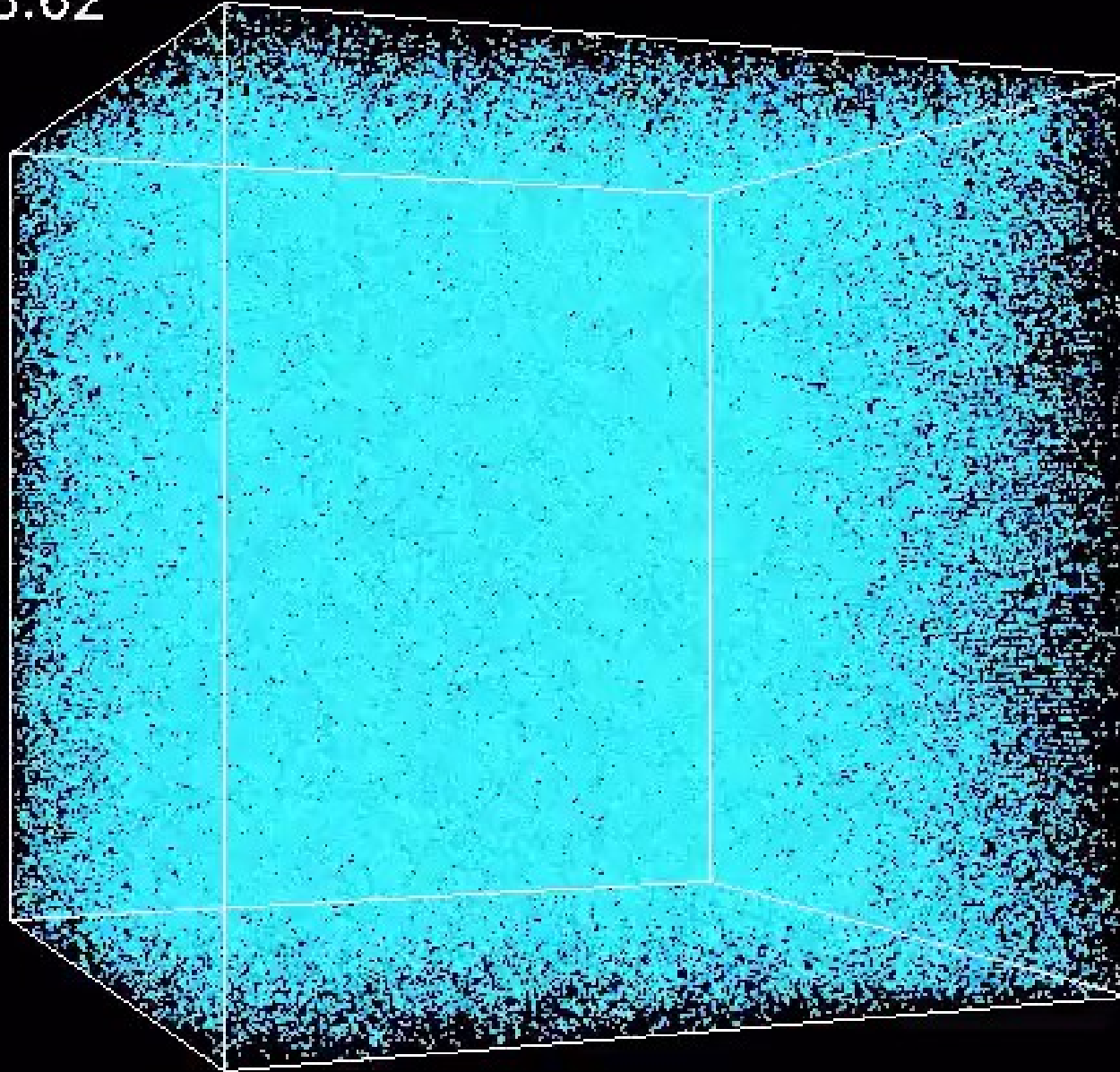
Introduction

- Galaxies are a building blocks of the Universe
- In the bottom-up model of galaxy formation, large mass galaxies are formed through mergers of smaller galaxies.



Credit: V. Springel

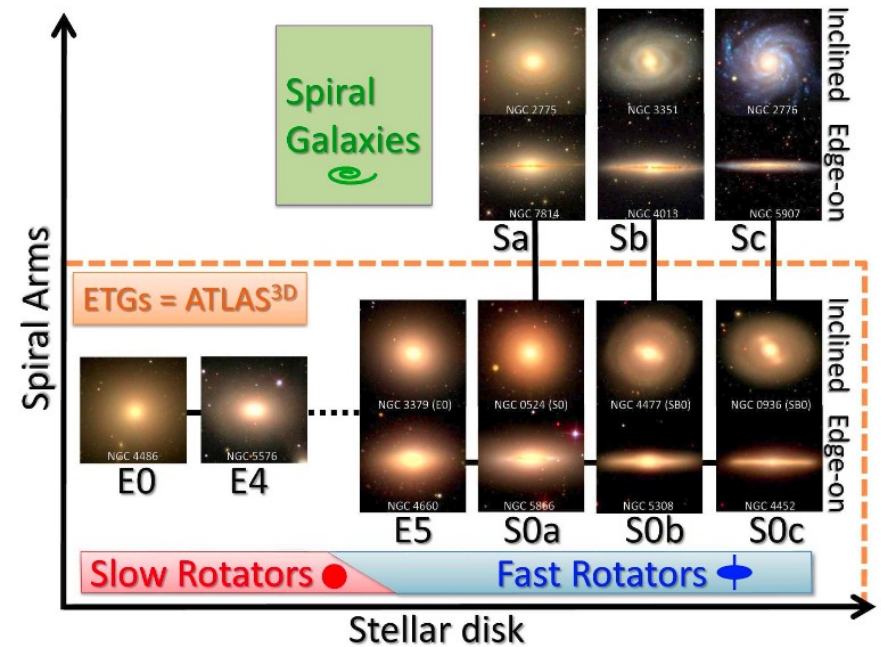
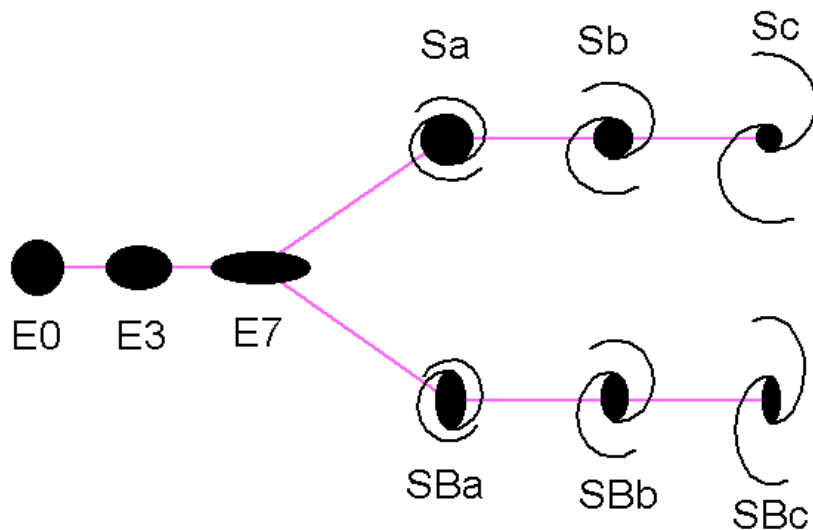
$Z=28.62$



Morphology clasification

Different morphologies,
masses, formation histories

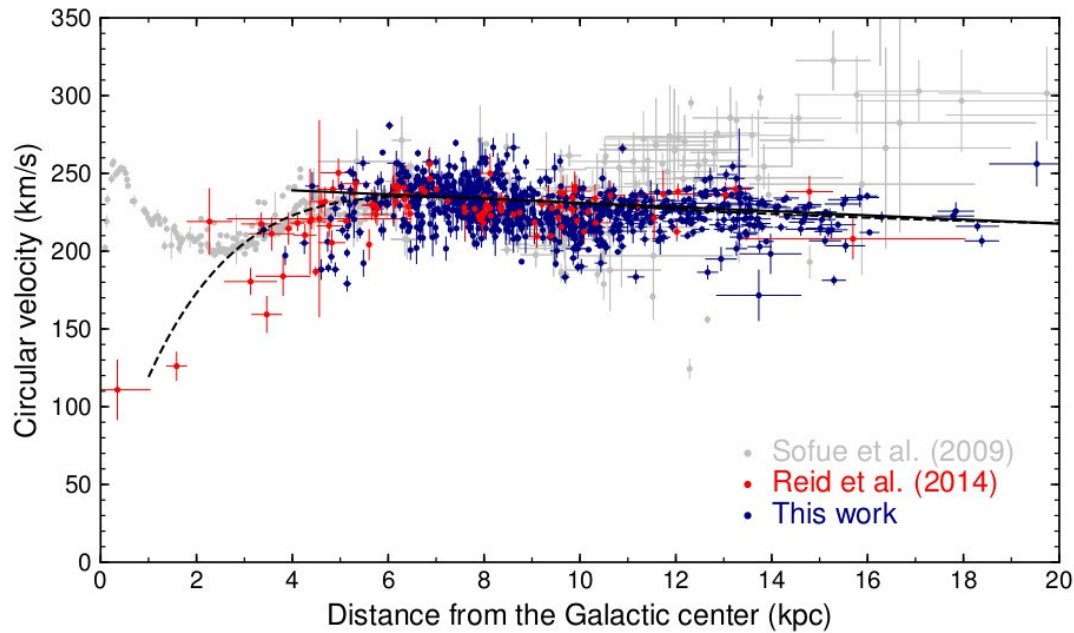
Hubble “Tuning Fork” Diagram



Cappellari et al. 2011

Spiral galaxies

- Disk stability (Toomre)
- Galactic rotation curves (Rubin)
- Dark matter in galaxy clusters (Zwicky)

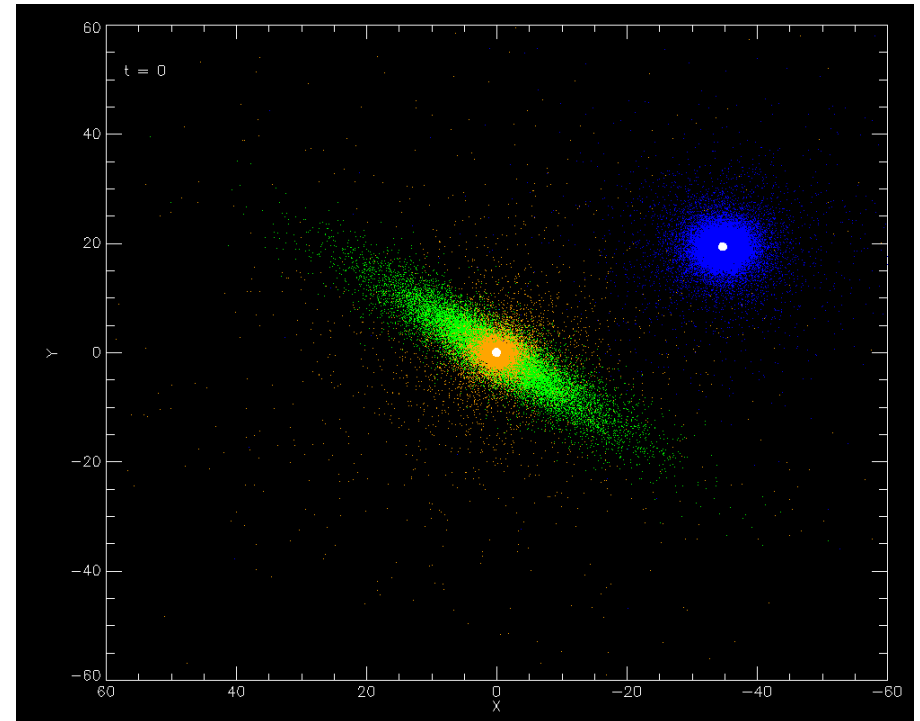
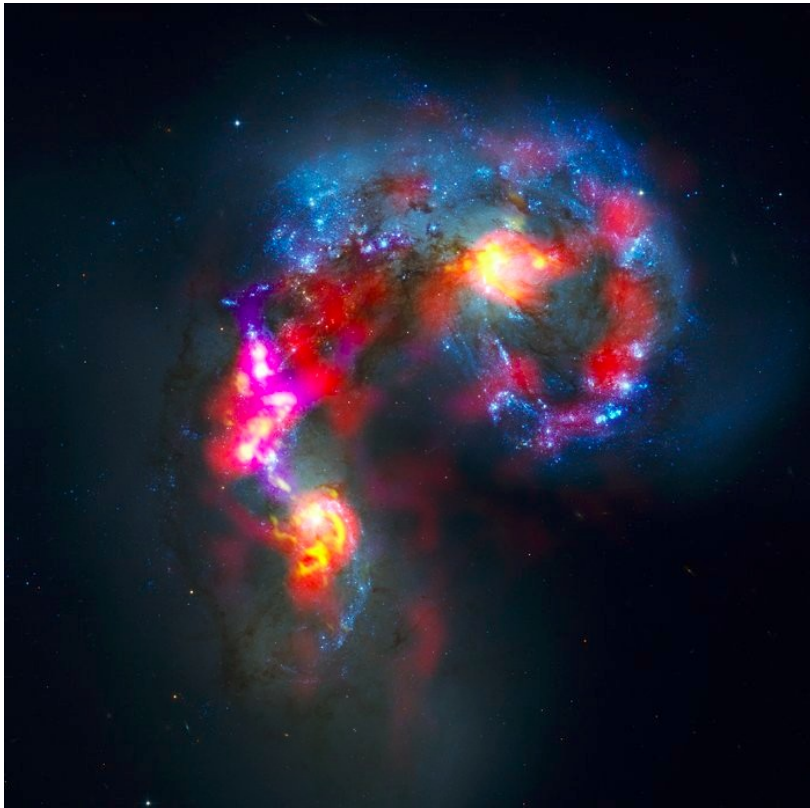


Mroz et al. 2019



Galaxy interactions

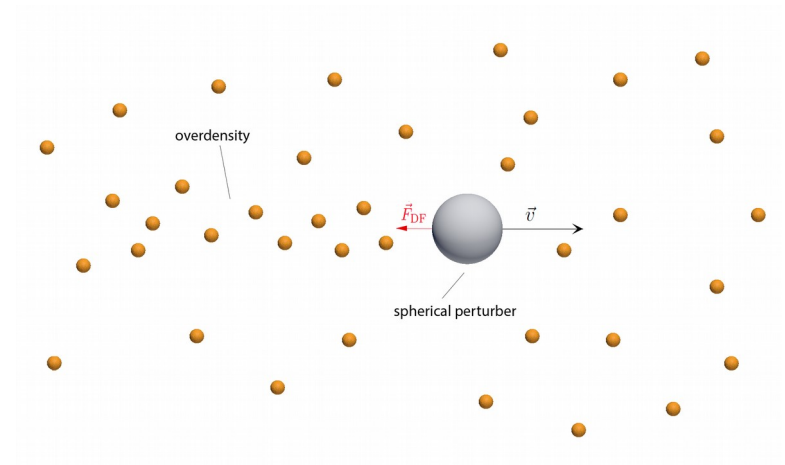
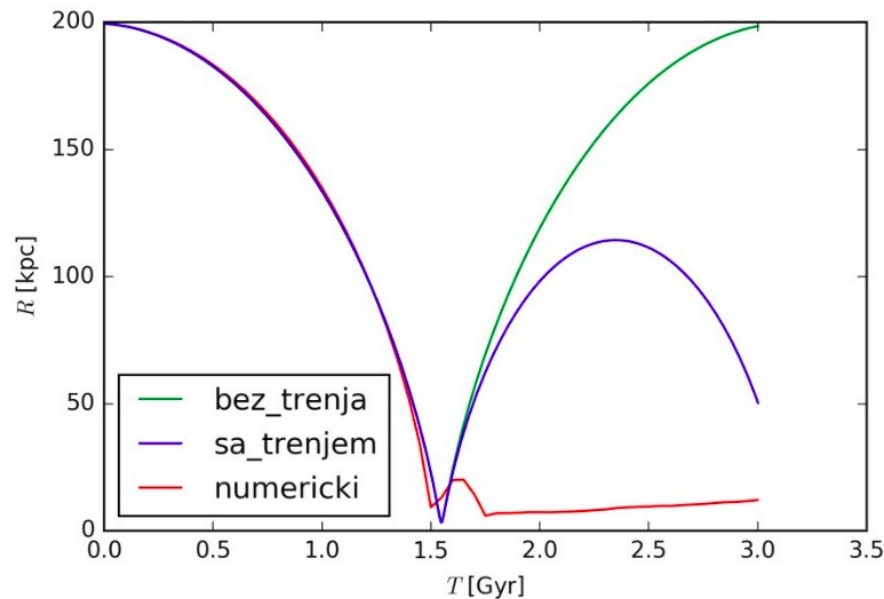
- Tidal disruption
- Tidal stripping
- Fly bys
- Mergers



Fardal et al. 2003

Dynamical friction

- Important for mergers
- Accounted in N-body models



Model of the spiral galaxy

(Geehan et al. 2006; Sadoun et al. 2014)

Disk

$$\Sigma(R) = \Sigma_0 e^{\frac{-R}{R_d}} = \frac{M_d}{2\pi R_d^2} e^{\frac{-R}{R_d}}$$

$$\propto \operatorname{sech}^2 \frac{z}{z_0}$$

$$\rho(R, z) = \frac{\Sigma(R)}{2z_0} \operatorname{sech}^2 \frac{z}{z_0}$$

Bulge

$$\rho_b = \frac{M_b r_b}{2\pi r (r + r_b)^3}$$

Halo (NFW, Navarro et. al 1996)

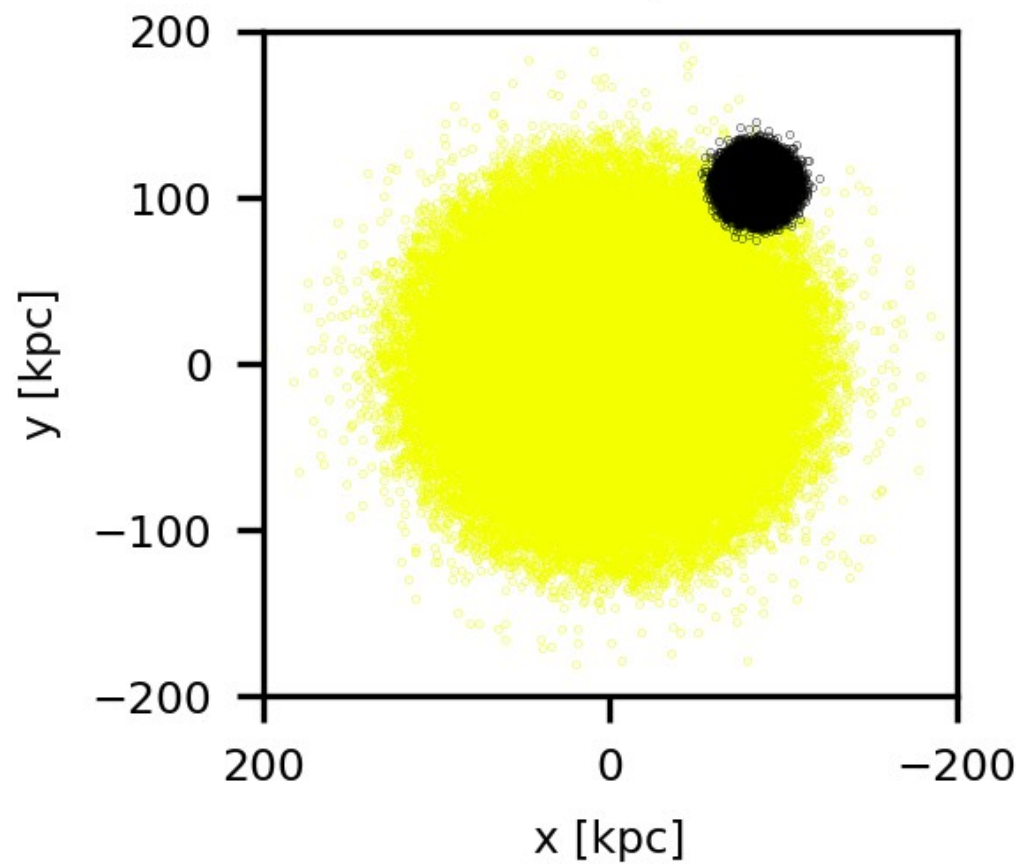
$$\rho_h(r) = \frac{\rho_0}{\frac{r}{r_h} (1 + \frac{r}{r_h})^2}$$

Stellar stream in M31

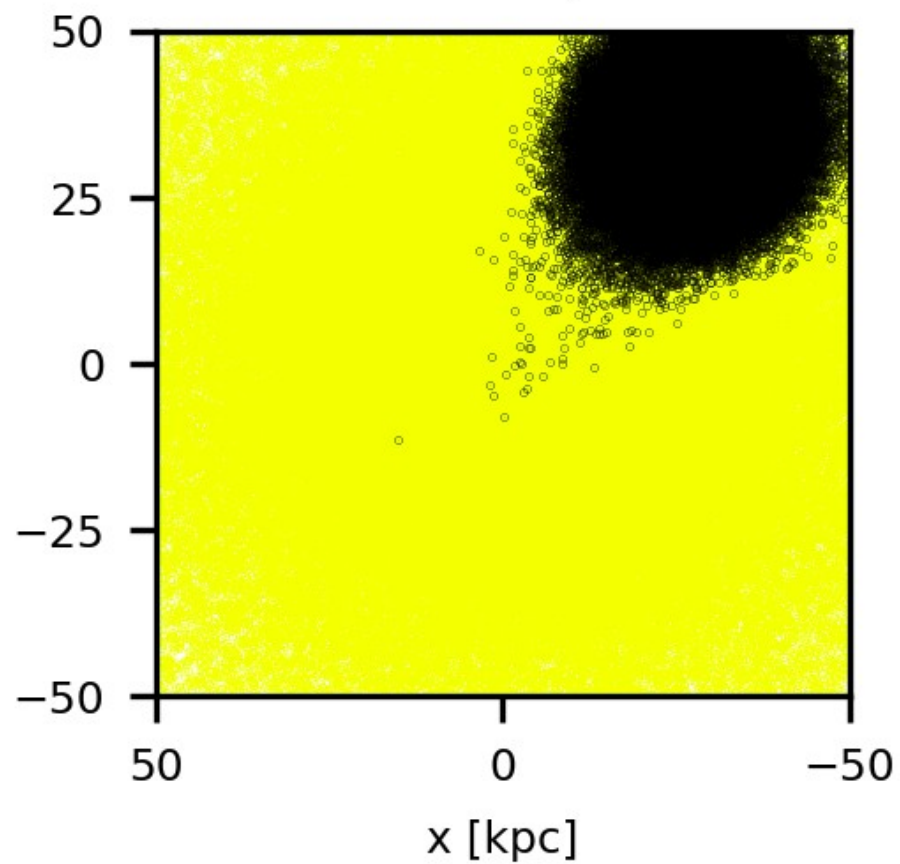
- The Giant Southern Stream is a faint stellar structure
 - Analytical simulation (Fardal et al. 2007)
 - N-body simulation (Sadoun et al. 2014; Micky & Mory 2016; Milošević et al. 2022)
 - Observations (Ibata et al. 2001; McConnachie et al. 2003, Ibata et al. 2004, Conn et al 2016; Cohen et al. 2018, Escala et al. 2020)
 - Timescale of the merger event
 - Physical properties
-
- Initial conditions were generated with GalactICS package (Widrow et al. 2008)
 - We ran simulations with Gadget 2 code (Springel 2005)

Dark matter halo

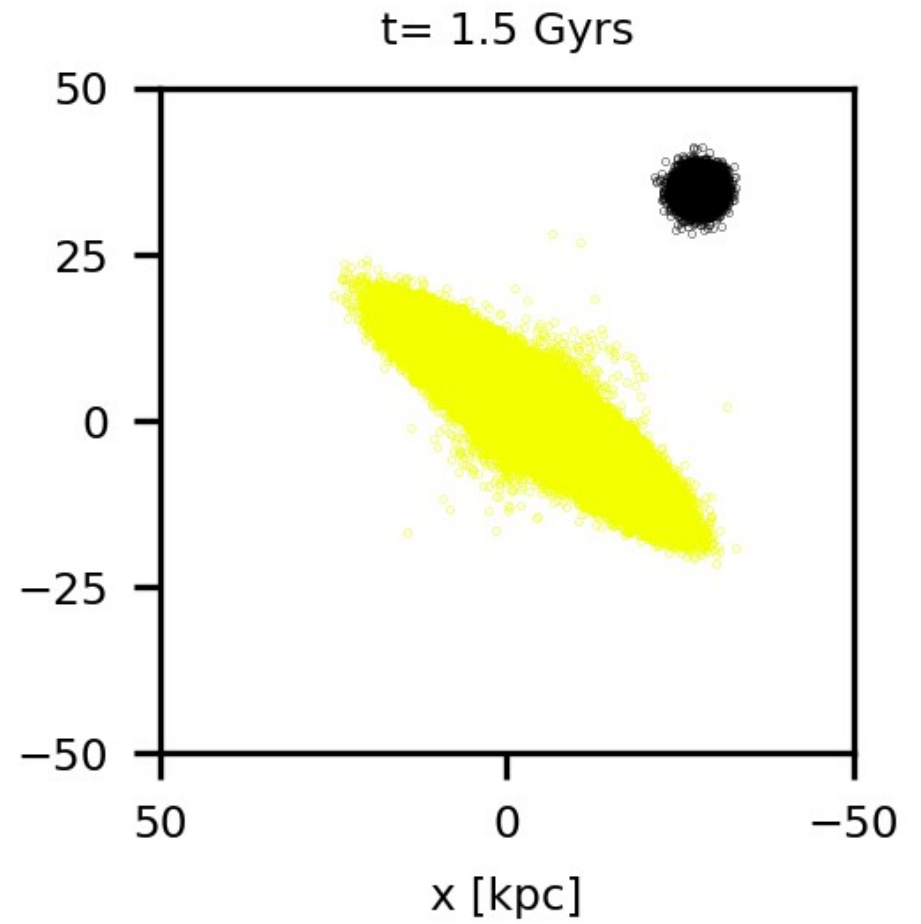
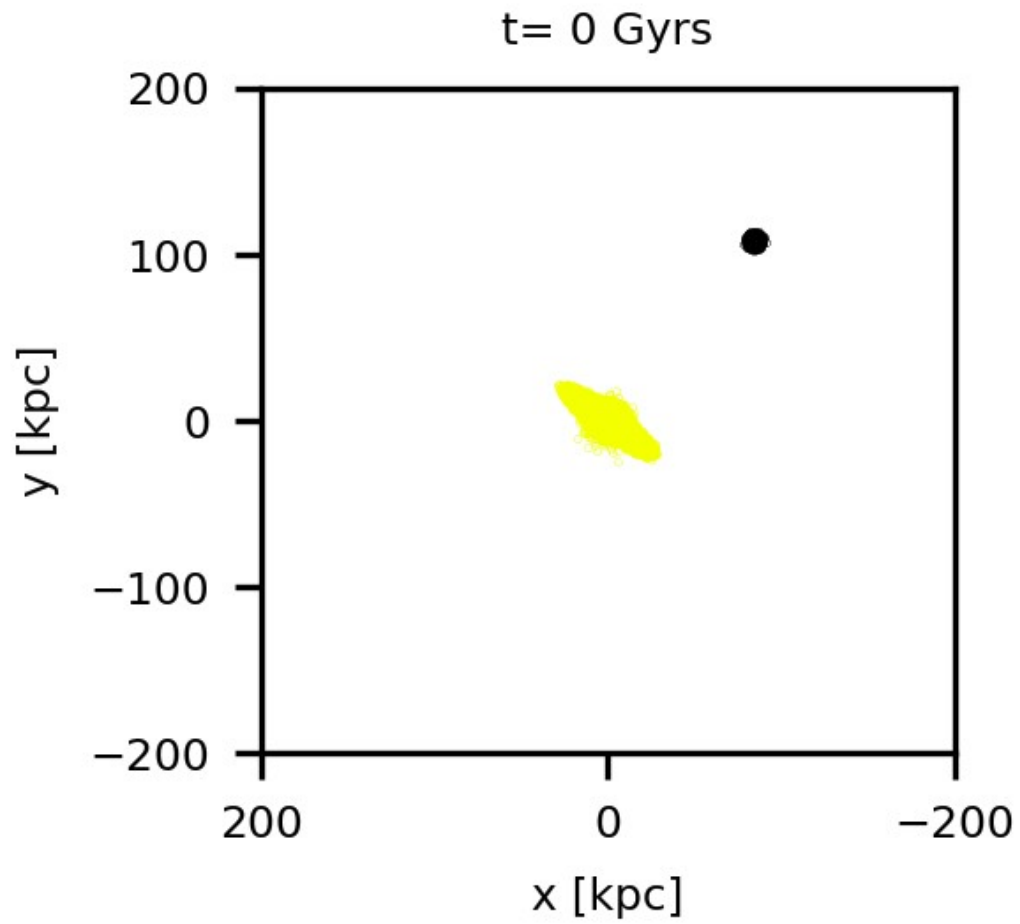
t= 0 Gyrs

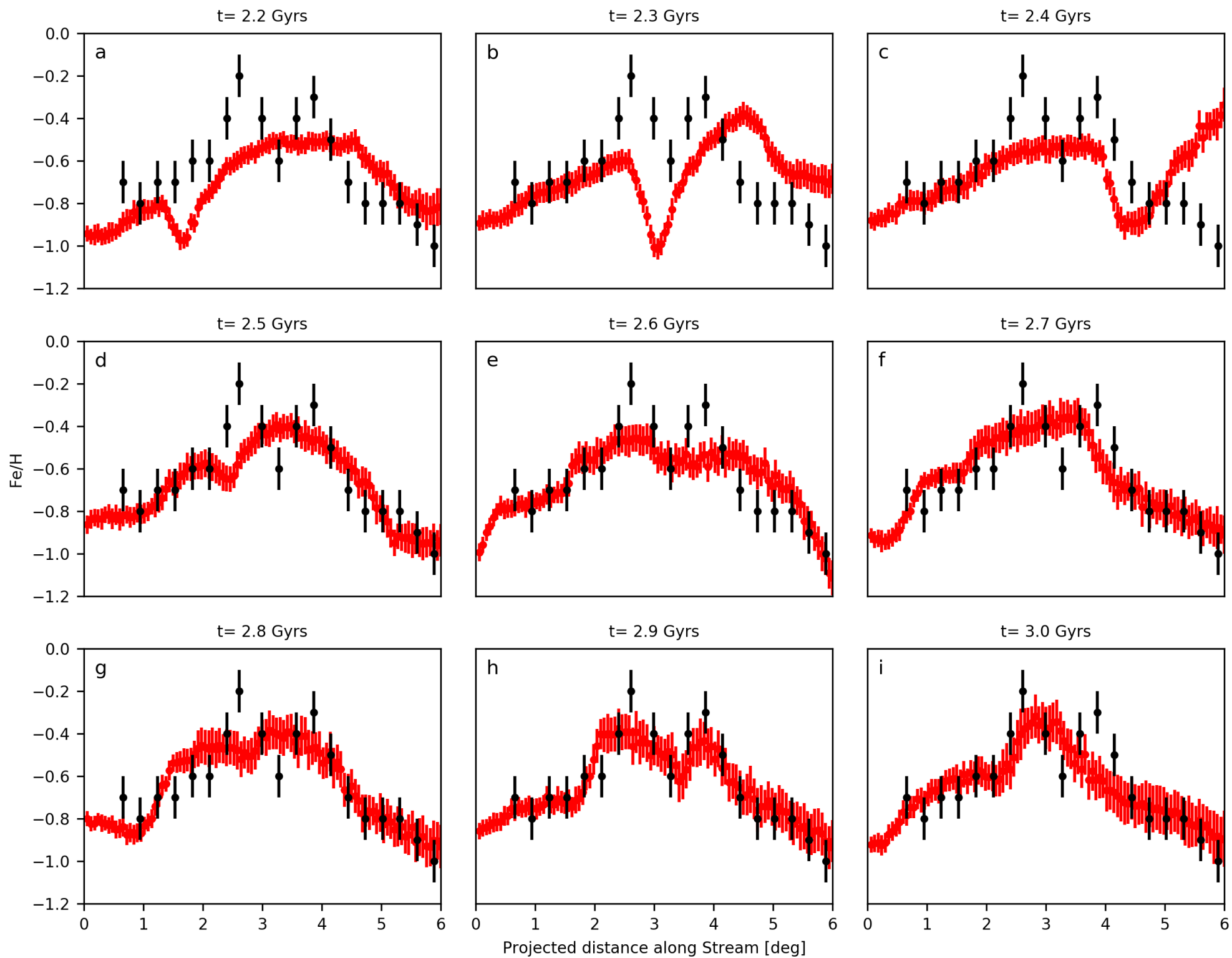


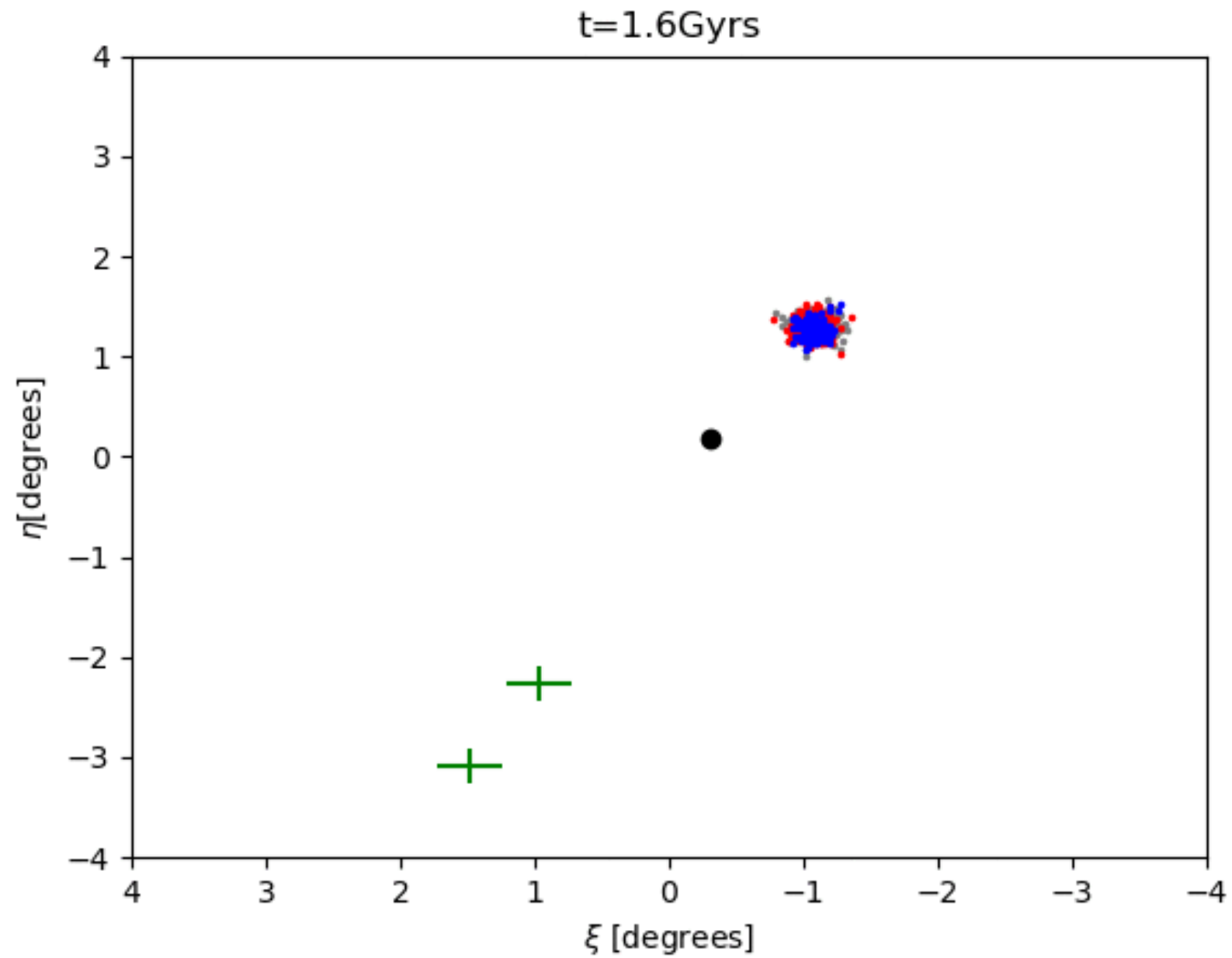
t= 1.5 Gyrs



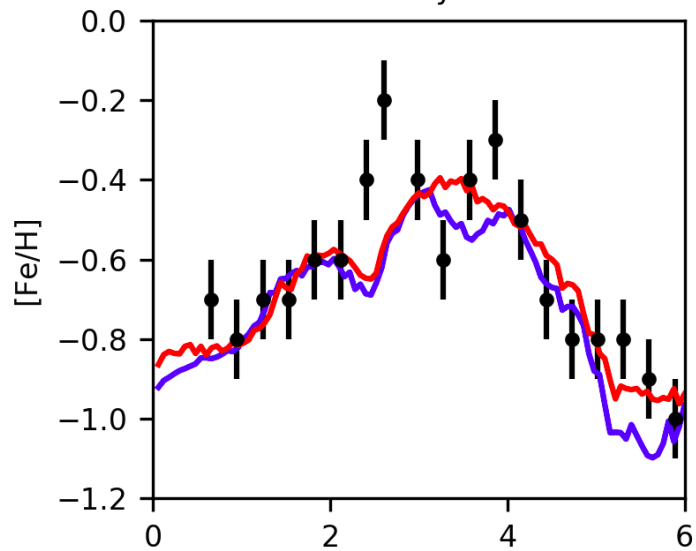
Baryonic matter



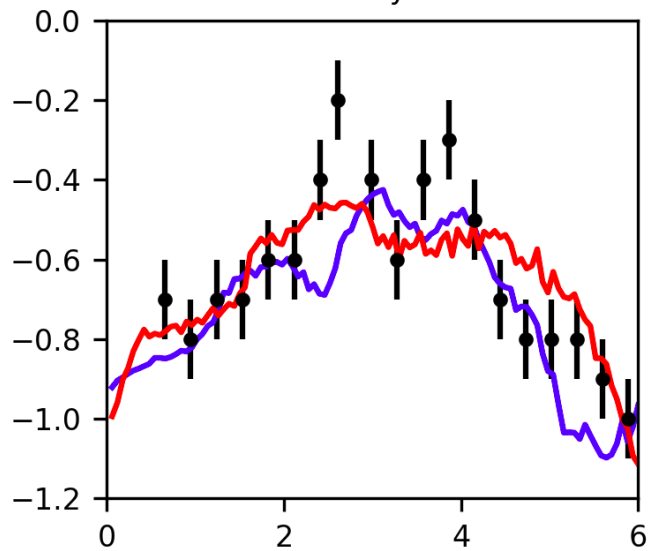




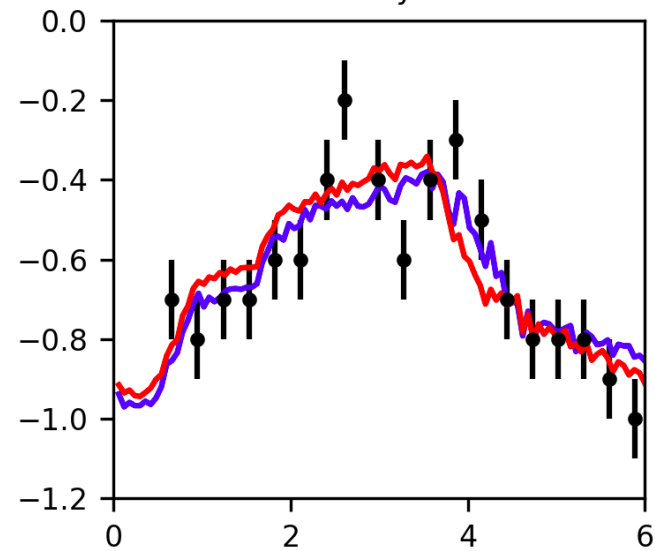
2.5 Gyrs



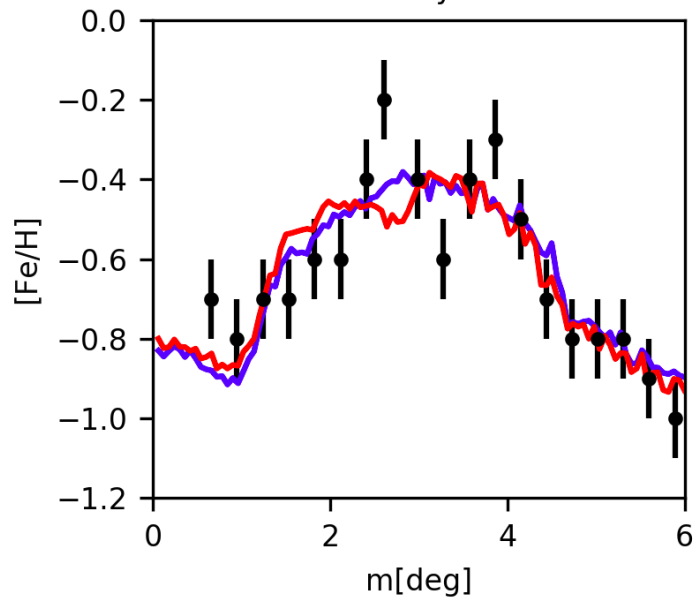
2.6 Gyrs



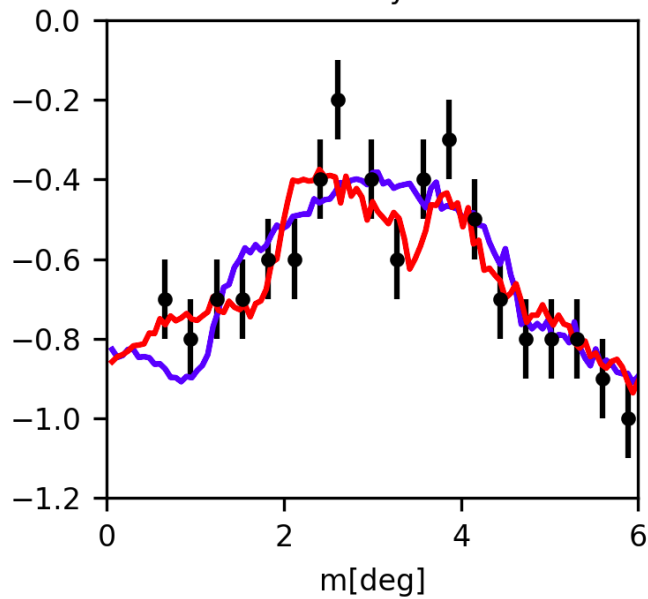
2.7 Gyrs



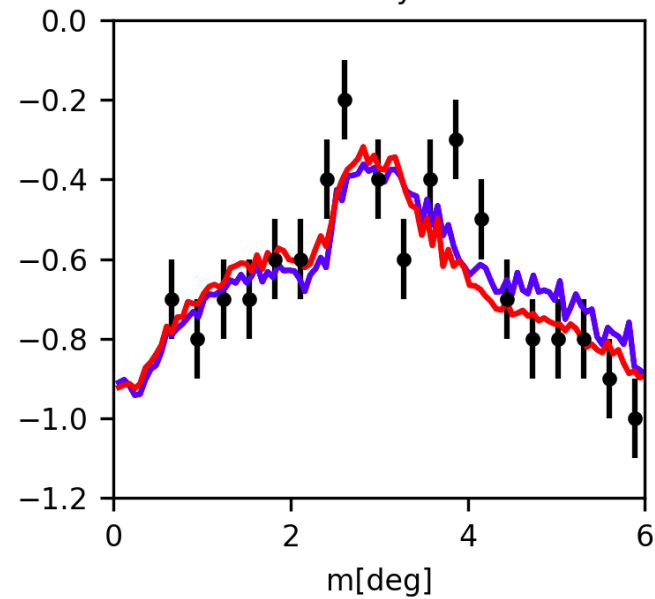
2.8 Gyrs

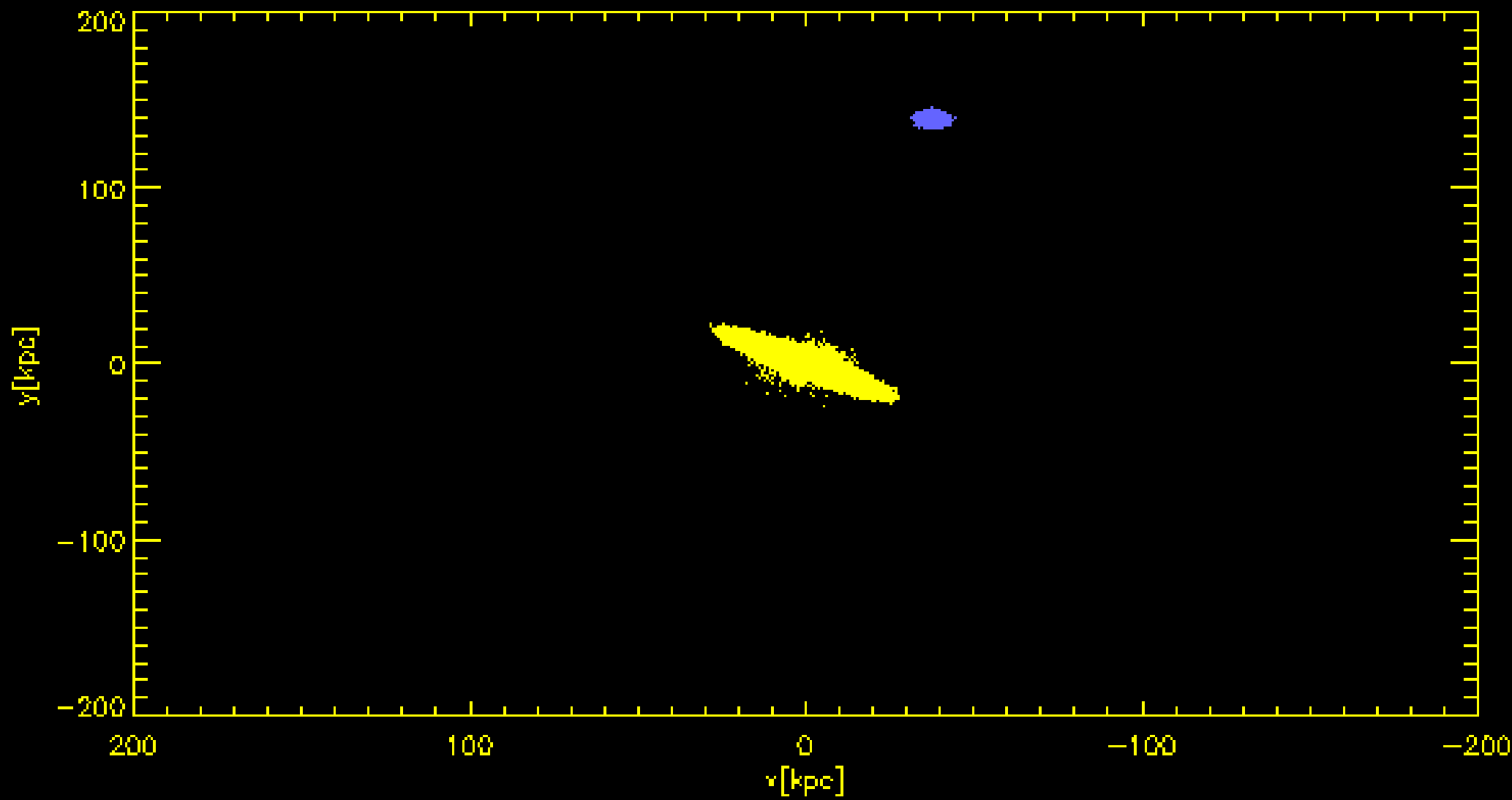


2.9 Gyrs

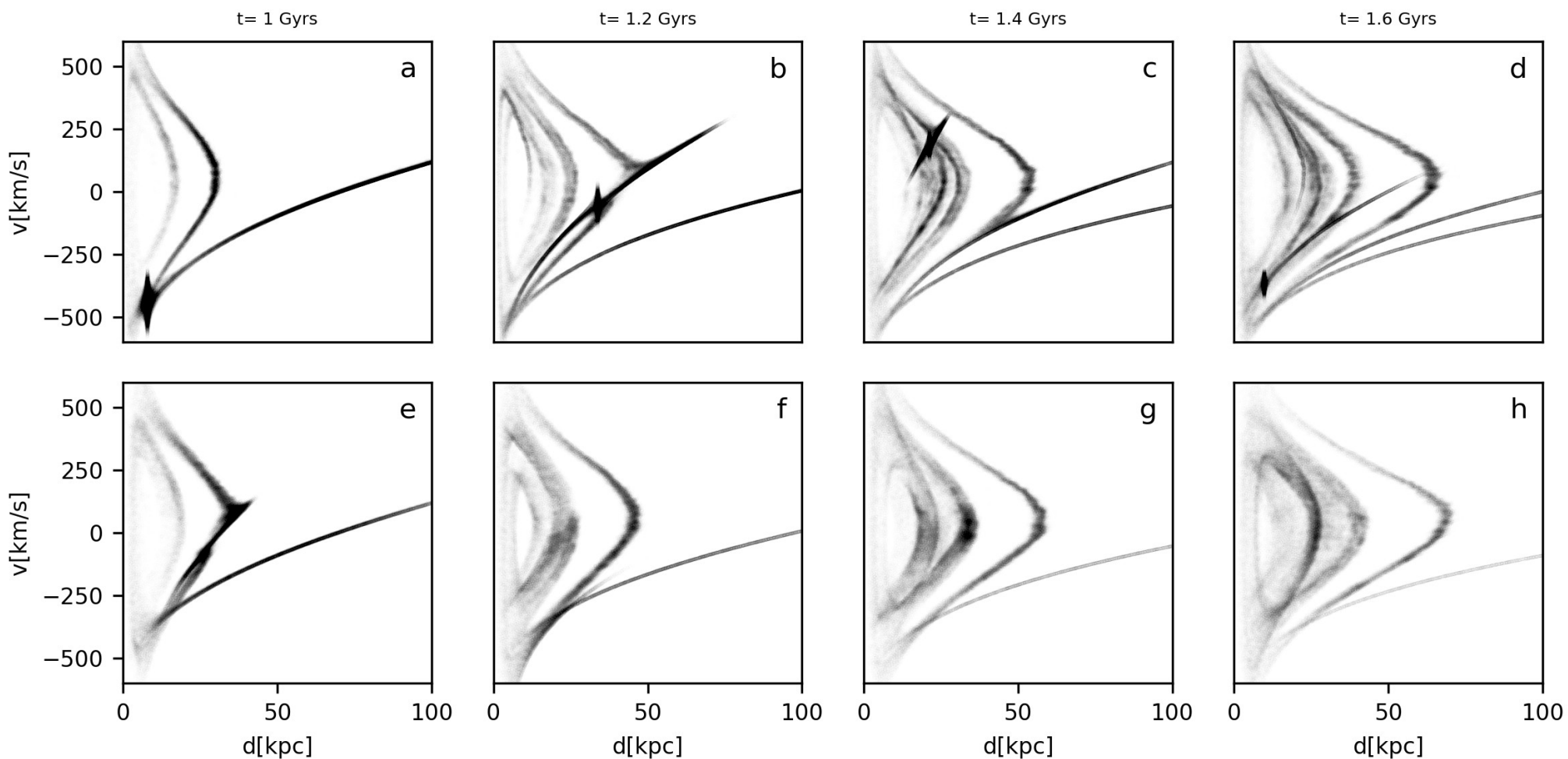


3.0 Gyrs





Simulations of shells in galaxies



Summary

- Density fluctuations and structure formation
- We use pure N-body simulation to describe dynamical history of galaxies
- Orbits of galaxies and tidal disruption
- Morphological substructures formation and timescale of events
- Mass of the dark matter halos in host and satellite galaxies
- Mass of the galaxies
- Future: gas in galaxies, other models for dark matter, and particle physics approach

Thank you!

