POWER-LAW TAILS OF THE DENSITY DISTRIBUTION IN STAR-FORMING CLOUDS: POSSIBLE EFFECTS OF ROTATION AND THERMODYNAMICS

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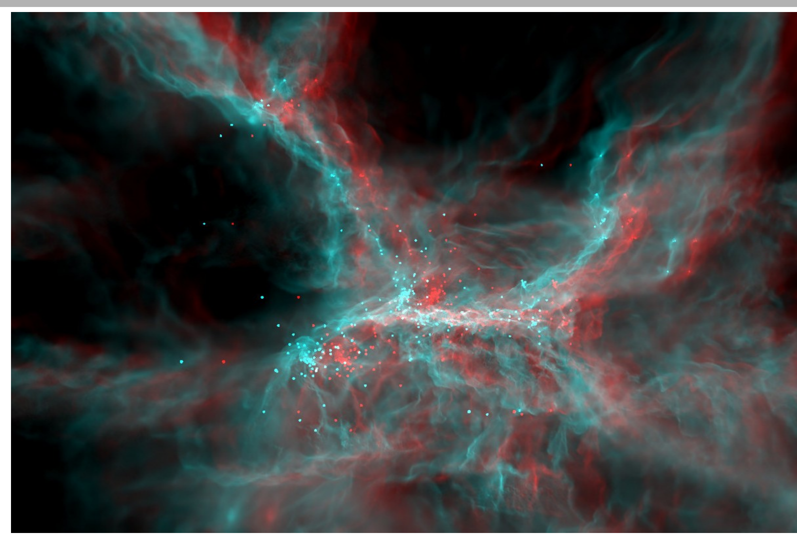
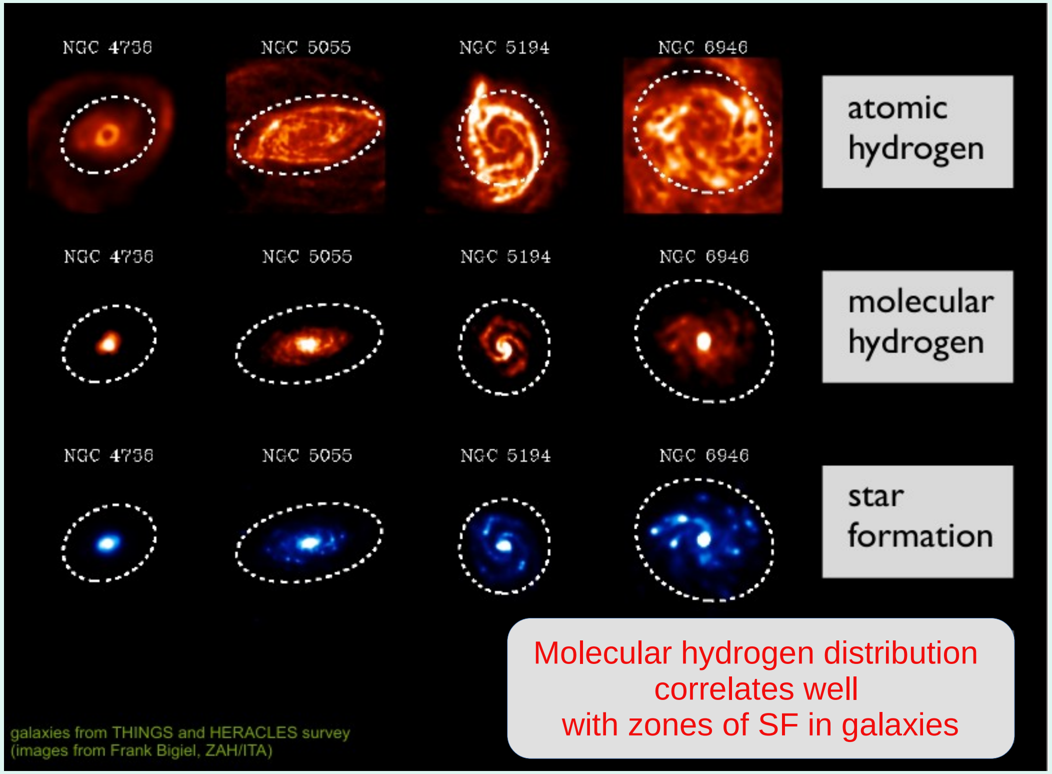


XIII[th] Bulgarian-Serbian Astronomical Conference
3-7 October 2022, Velingrad, Bulgaiia

Molecular clouds (MCs) as sites of star formation (SF)



Molecular hydrogen distribution correlates well with zones of SF in galaxies

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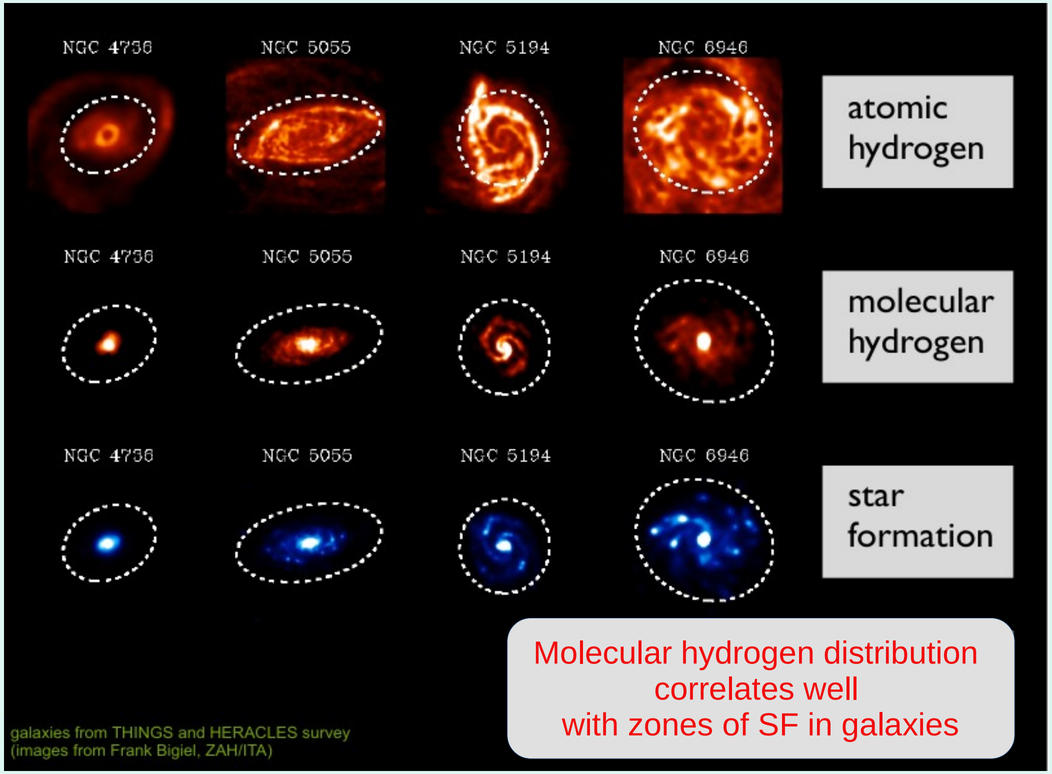

Tracers of molecular gas

Dust extinction
('dark clouds')

MC ρ Ophiuchi
(image: HST, DSS1; sky.esa.int)

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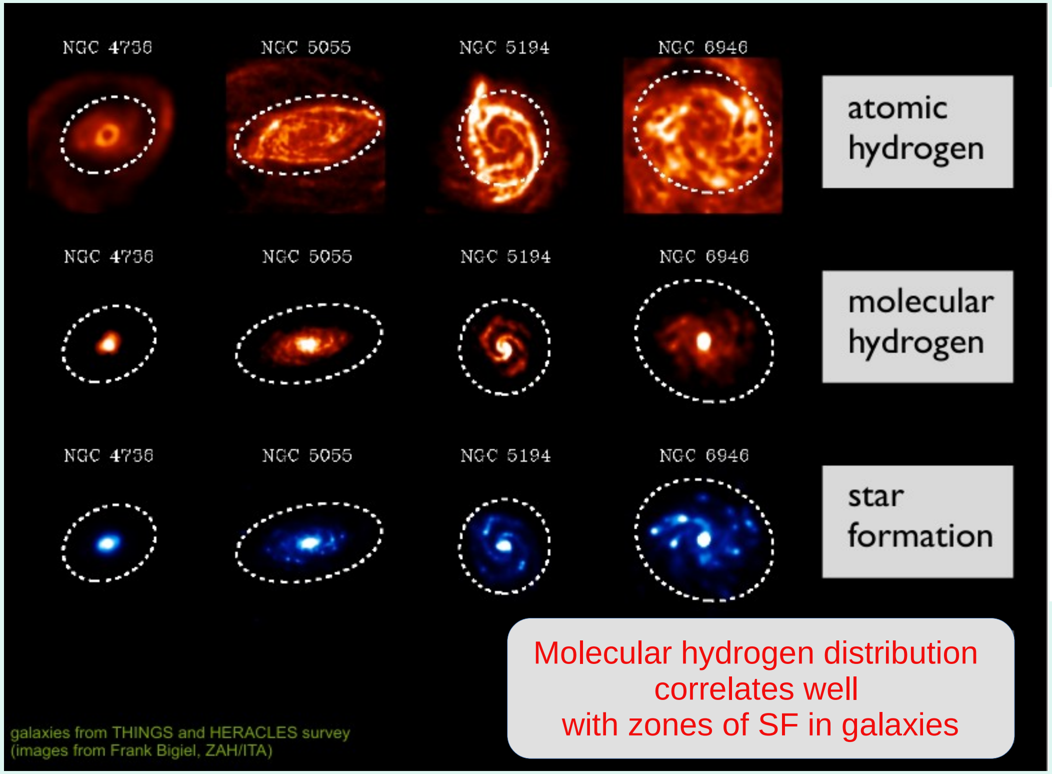

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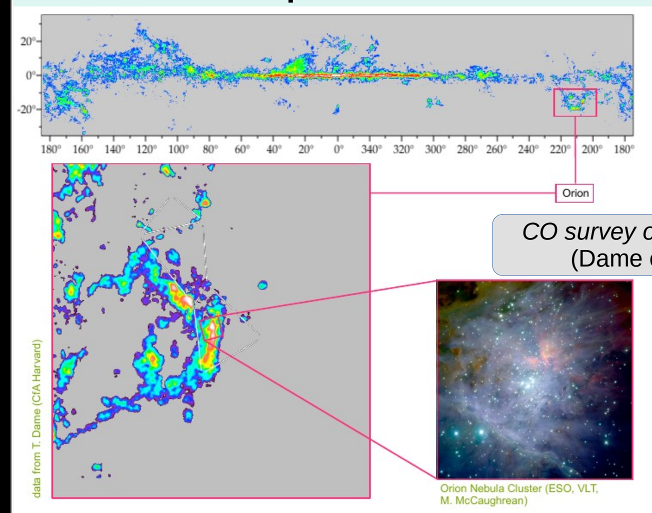
Tracers of molecular gas

Emission of CO species



CO survey of the Milky Way
(Dame et al. 2001)

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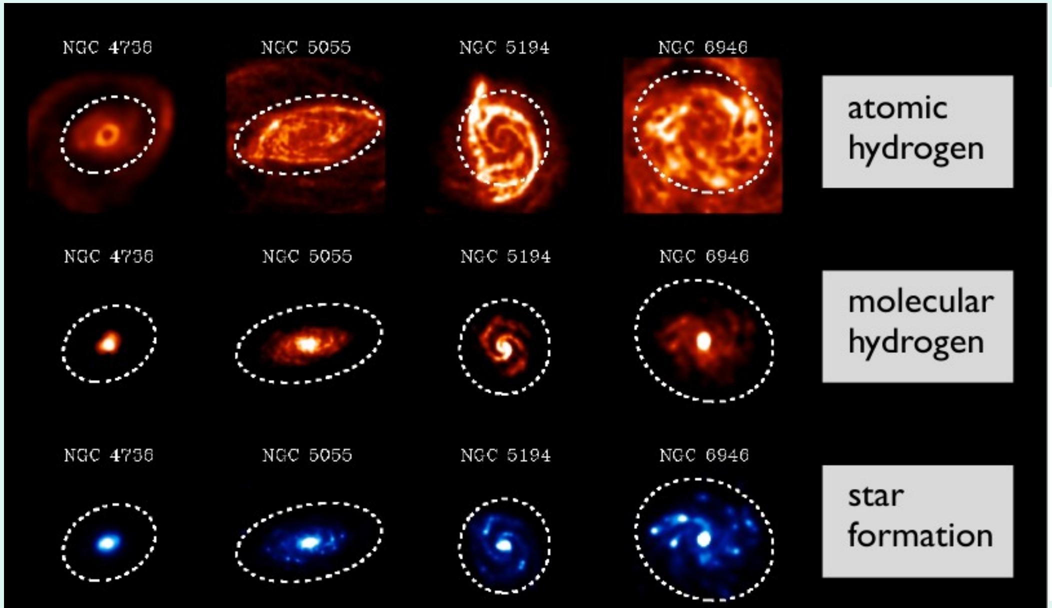

Molecular hydrogen distribution correlates well with zones of SF in galaxies

galaxies from THINGS and HERACLES survey
(images from Frank Bigiel, ZAH/ITA)

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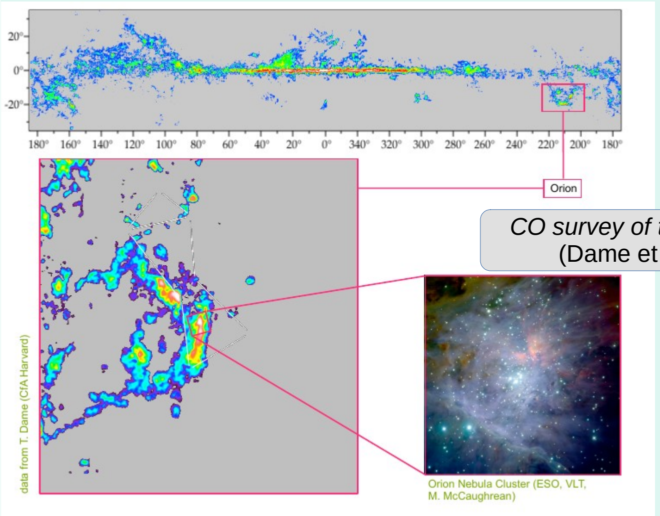

CO survey of the Milky Way
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Dust emission
(small-scale structure)

MC ρ Ophiuchi
(image: Herschel SPIRE 250, 350, 500 µm)

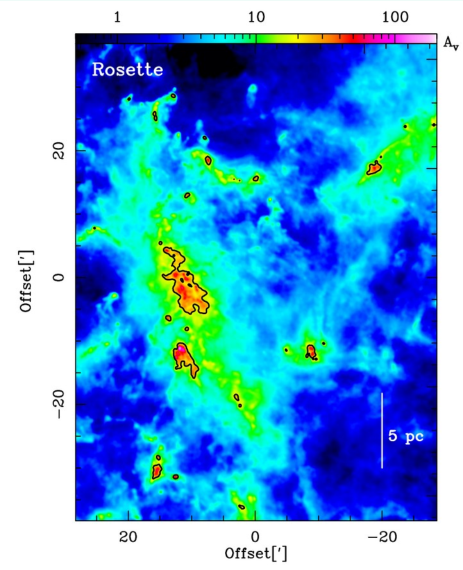
Dust extinction
('dark clouds')

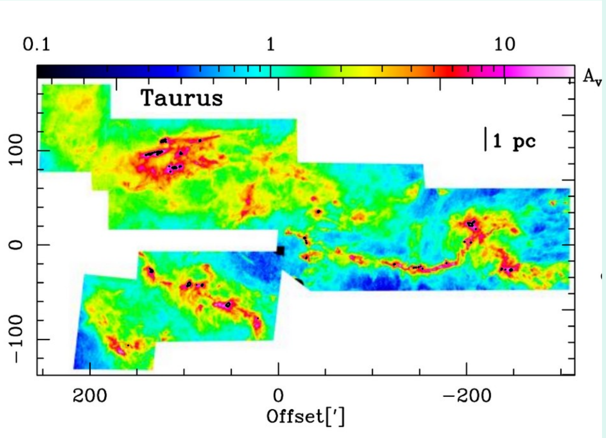
MC ρ Ophiuchi
(image: HST, DSS1; sky.esa.int)

Appropriate
for study of star-forming clouds and their substructures (e.g., pre-/protostellar cores)

The variety of star-forming activity in molecular clouds (MCs)

Herschel imaging at high angular resolution (18 arcsec; Schneider et al. 2022)





High-mass SF clouds

- giant MCs
- sizes: up to 100 pc
- masses: 10^5-10^6 $M_\odot$
- Signatures of high-mass and cluster formation, massive, grav. unstable filaments of high column-density

Low-mass SF clouds

- sizes: up to 10-30 pc
- masses: 10^3-10^4 $M_\odot$
- They form typically low-mass stars

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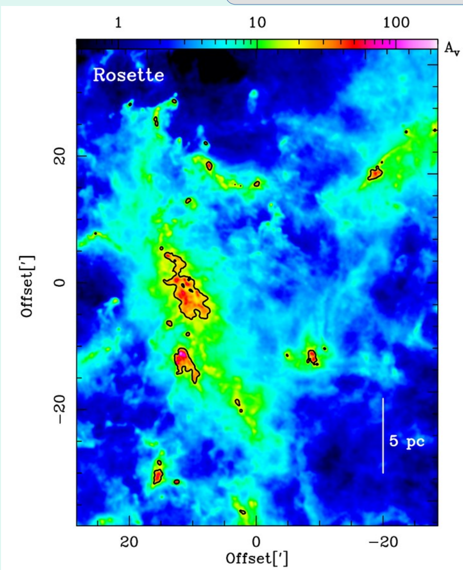

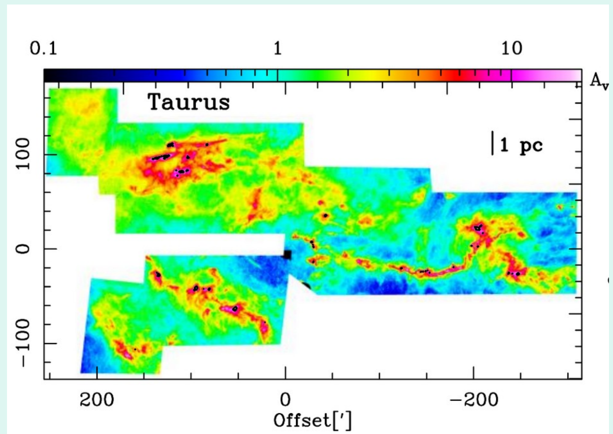

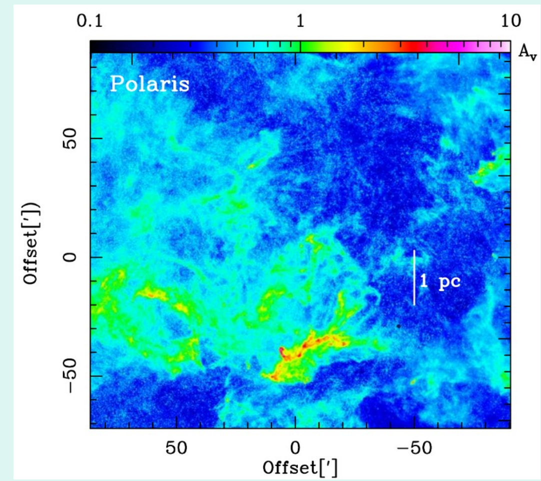

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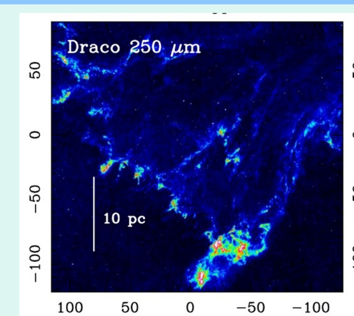
Quiescent clouds

- Poor or no SF activity



Diffuse clouds

- Mostly atomic

The complex physics of star-forming MCs

- The complex physics of MCs is governed by gravity, supersonic turbulence, magnetic fields and – in the general case – an isothermal equation of state (EOS).
- Accretion from the surrounding medium and feedback from new-born stars and supernovae play an essential role in cloud's evolution.
- Effects of rotation

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- Effects of rotation

This complex physics is imprinted in:

- General structure of MCs in terms of scaling relations of velocity dispersion and mass.
- Probability distribution of different quantities of the medium
- Physical parameters of substructures (clumps, cores, filaments)

(Column-)Density distribution as a research tool

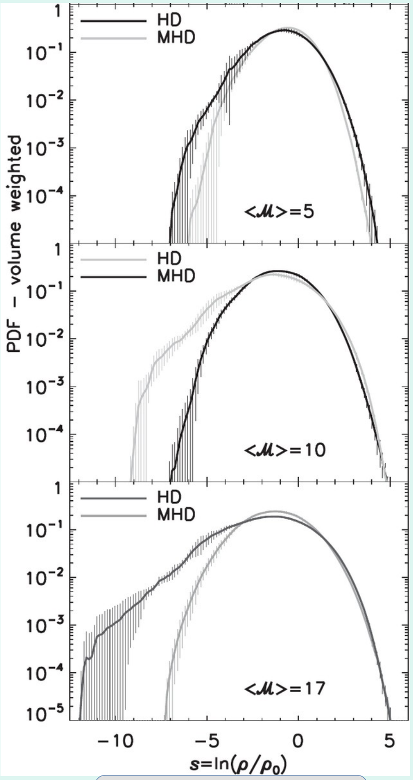
$s = \log(\rho / \langle \rho \rangle)$ $p_s \mathrm{d}s$ - probability distribution function (PDF) of logdensity

Lognormal (part of) PDF
→ isothermal supersonic turbulence

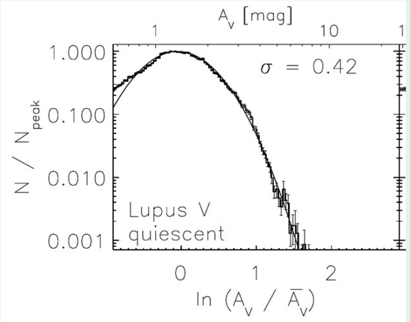
$$p_s \mathrm{d}s = \frac{1}{\sqrt{2\pi\sigma_s^2}} \exp\left[-\frac{(s - s_0)^2}{2\sigma_s^2} \right] \mathrm{d}s \qquad \sigma_s^2 = \ln\left[1 + b^2 \mathcal{M}^2\right]$$

simulations

observations



Molina et al. (2012)



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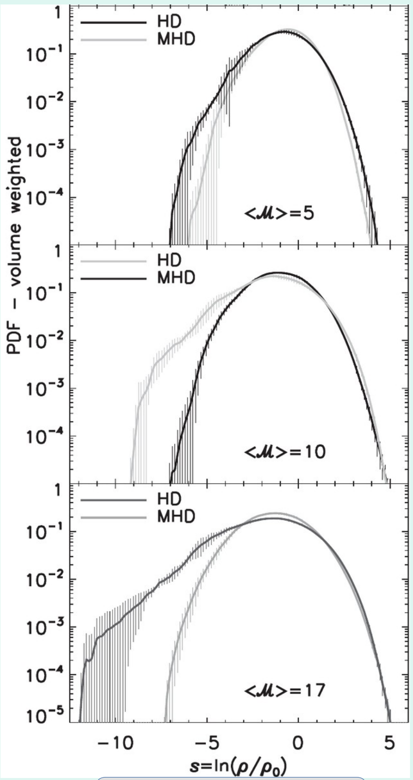

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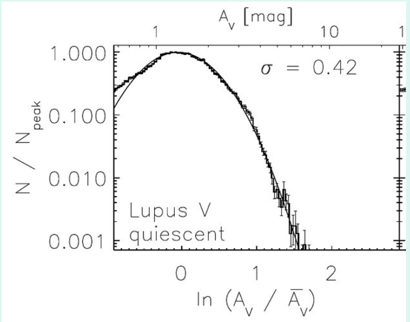
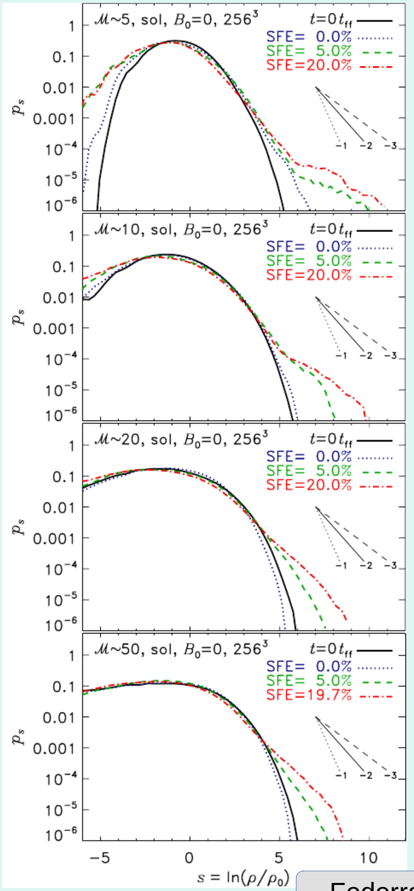


Kainulainen et al. (2009)

Emergence of a power-law tail (PLT)
→ increasing role of self-gravity

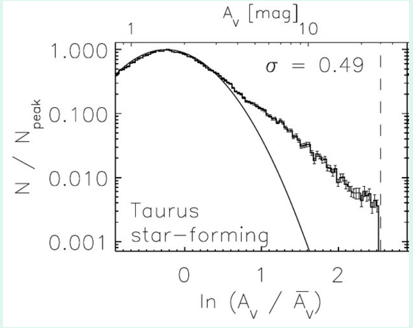
$$\mathrm{PLT} \propto \exp(qs), \quad q < 0$$

simulations



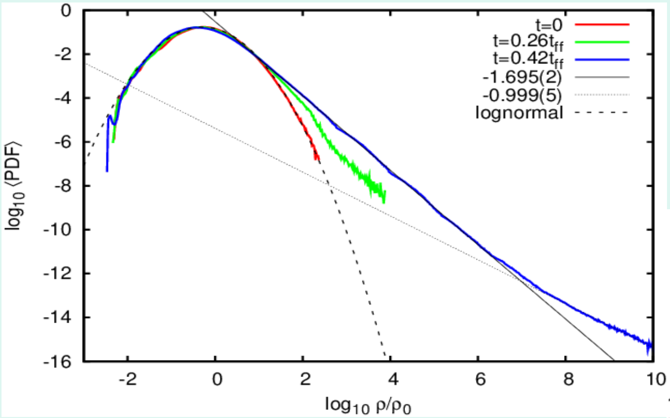
Federrath & Klessen (2013)

observations



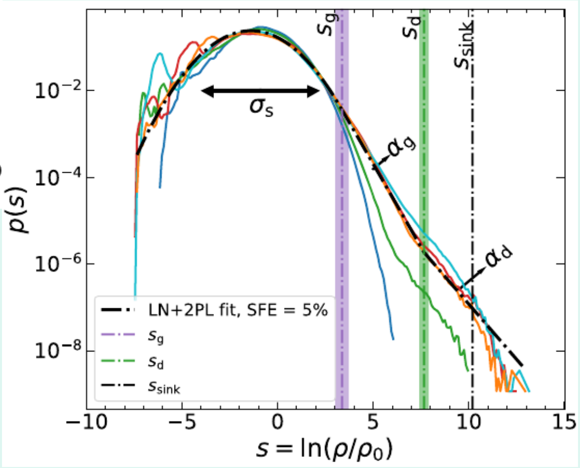
Kainulainen et al. (2009)

PDF of mass density (ρ-PDF) in evolved star-forming MCs



Kritsuk, Norman & Wagner (2011)

- HD simulations of supersonic, isothermal and self-gravitating turbulent medium.
- Resolution: down to AU scales in the dense cores.



Khullar et al. (2021)

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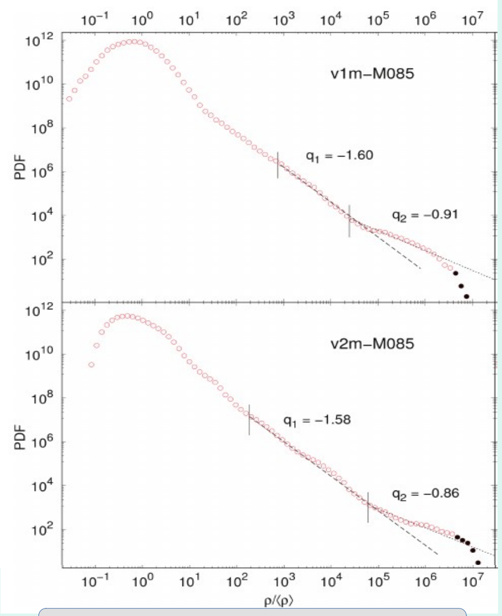

Marinkova et al. (2021)

- HD simulations of typical large SF clumps (0.5 pc), with large Jeans content (32, 354 M_J); variation of turbulent driving
- Resolution: down to ~3 AU.

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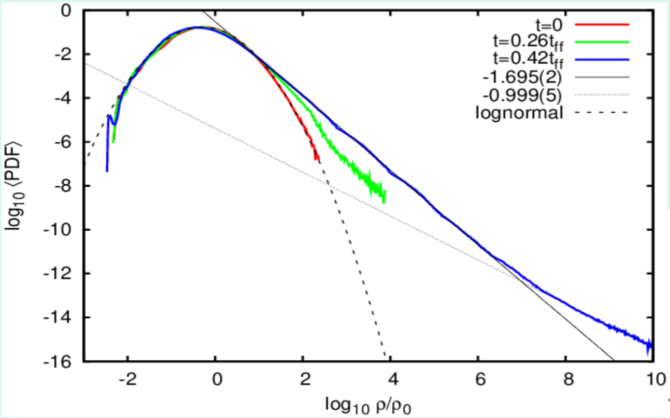

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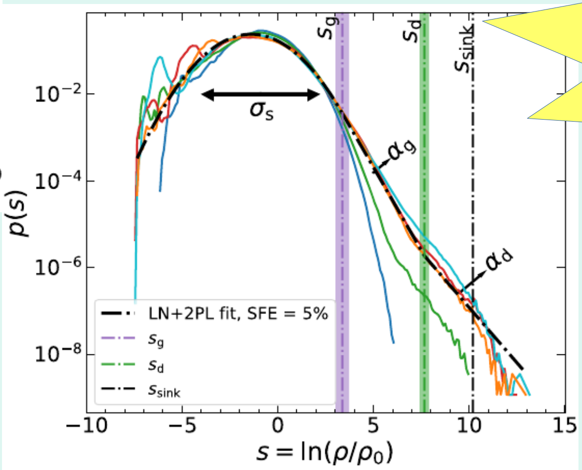

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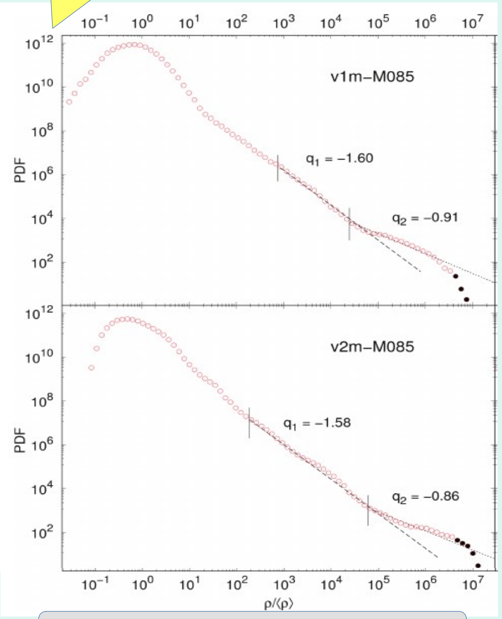
Emergence of a ***second*** *PLT at very high densities and at later evolutionary stages*



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N-PDFs of variety of MCs with various SF activity

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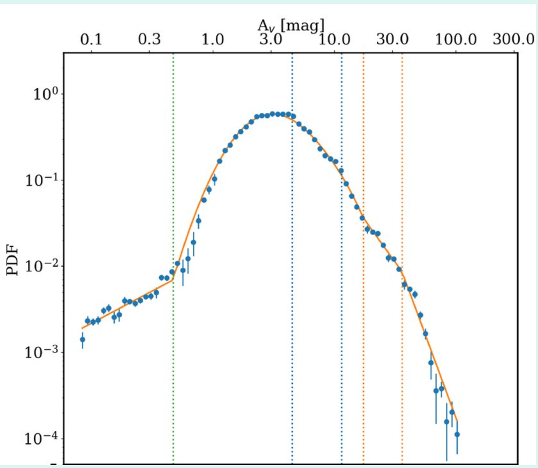
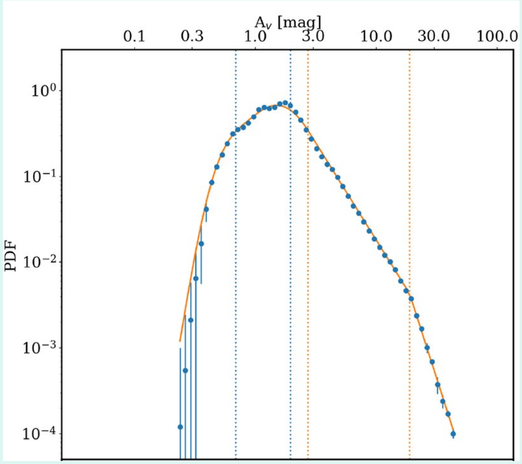
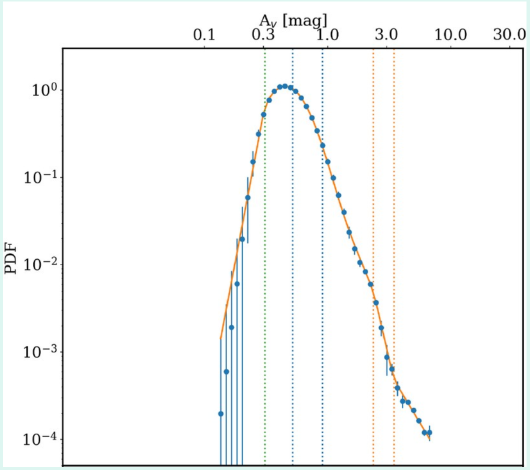
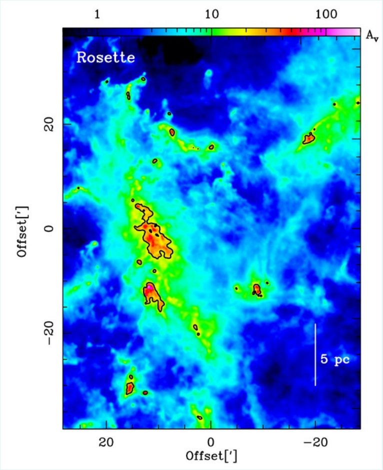
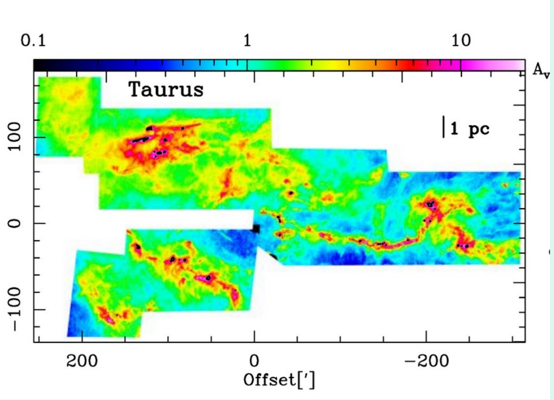
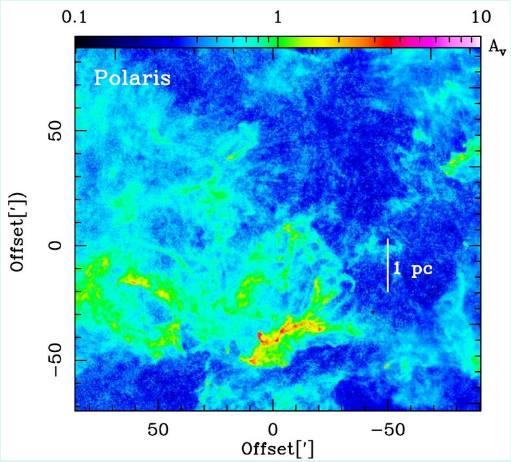












High-mass SF clouds

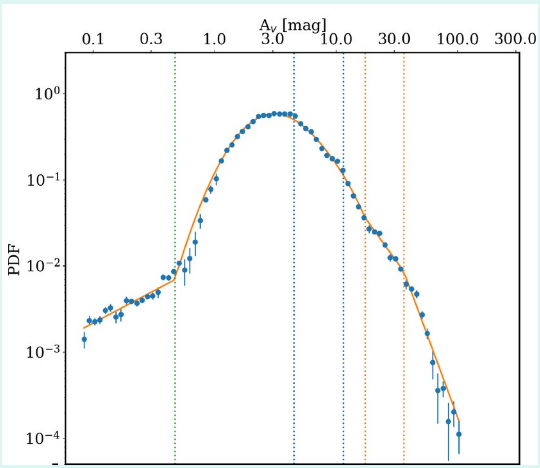
Low-mass SF clouds

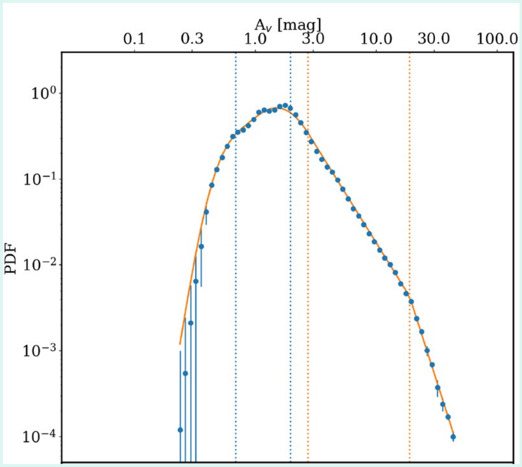
Quiescent clouds

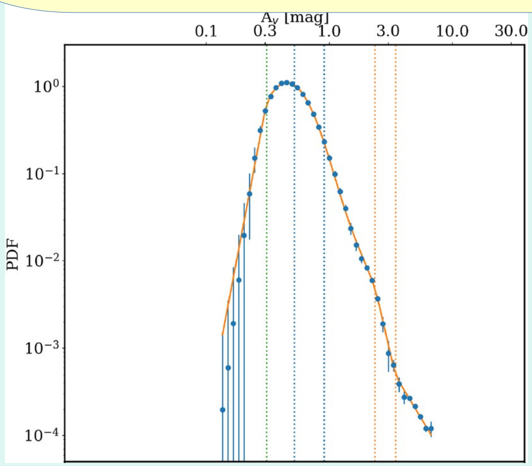
N-PDFs of variety of MCs with various SF activity

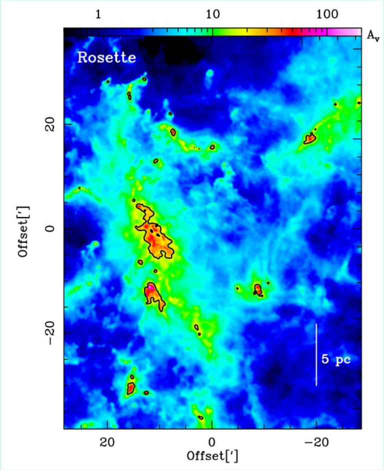
Herschel imaging at high angular resolution (18 arcsec; Schneider et al. 2022)

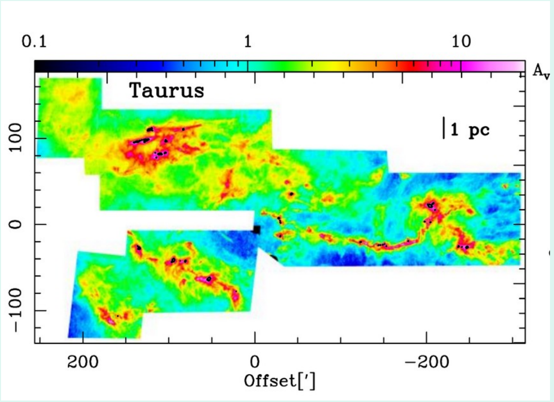
*Double PLTs **confirmed** from observations!*

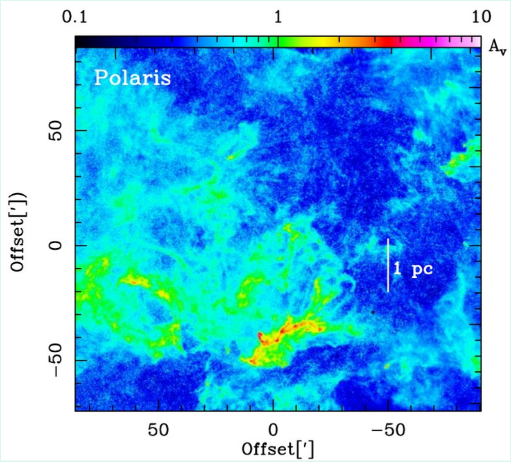










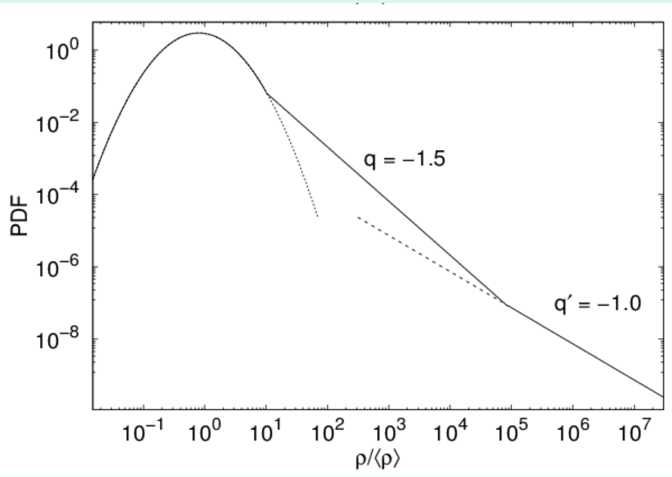


High-mass SF clouds | **Low-mass SF clouds** | **Quiescent clouds**

Some suggested explanations of the second PLT

- Rotation of prestellar cores (Kritsuk et al. 2011), structures in rotationally flattened disks (Murray et al. 2017)
- Changing balance between gravity and turbulence in the course of MC evolution: first PLT signifies (Murray et al. 2017)
- Amplification of magnetic fields in the densest clumps within the cloud (Schneider et al. 2015)
- Change in thermodynamics: transition from isothermal state (at larger scales) to polytropic state (at small scales) in self-gravitating clouds with steady-state accretion (Donkov et al. 2021)

All those factors act together?



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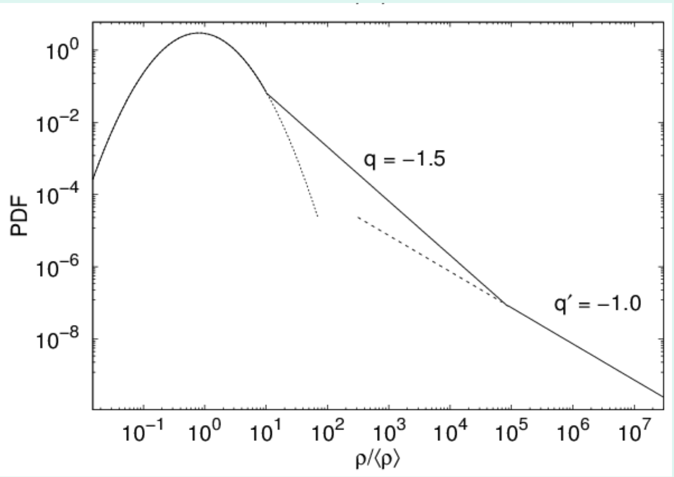
This report: study of ρ-/N-PDF evolution which allows to distinguish the effect of rotation on the high-density end

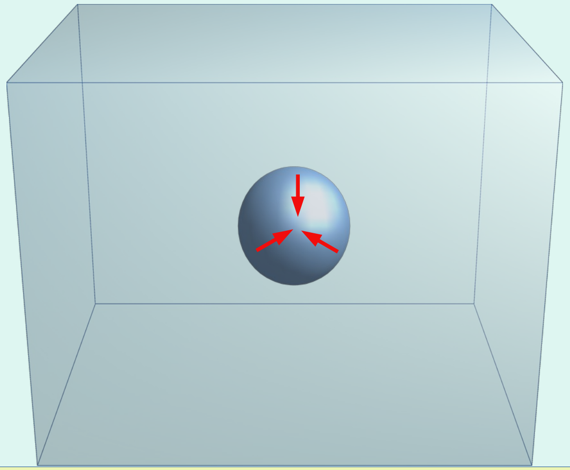


Used data and applied method

Numerical simulations (Wollenberg et al. 2020)

- Voronoi moving-mesh code AREPO (Springel 2010).
- Simulated contracting SF clump: single Bonnor-Ebert sphere within a homogeneous 13 pc box.
- Primordial gas: a network of 45 chemical reactions between different species of H and He and free electrons provides for treatment of cooling and for computation of the polytropic index
- Different physical setups
  - Pure infall (PI)
  - Rotation only (RO), β=0.01 and β=0.10
  - Turbulence only (TO), α=0.05 and α=0.25
- Run times: ~2 τ_{ff}; number of protostars formed: from 1 (PI) up to a few dozens

α Turbulent vs. gravitational potential energy ratio
β Rotational kinetic vs. gravitational potential energy ratio



'contracting SF clump'
BE-sphere: R~2 pc, 2.6x10^3 $M_\odot$

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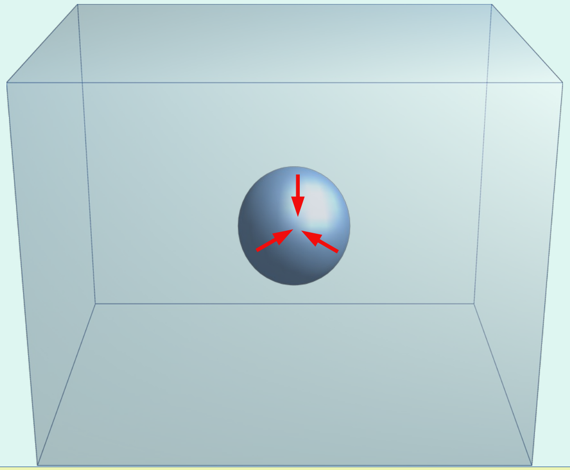

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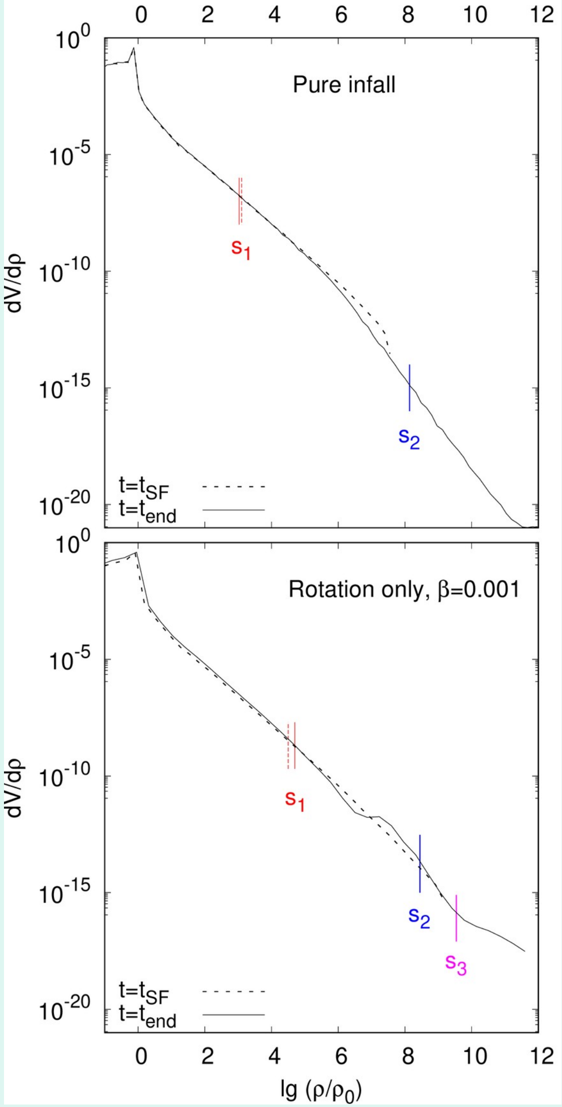
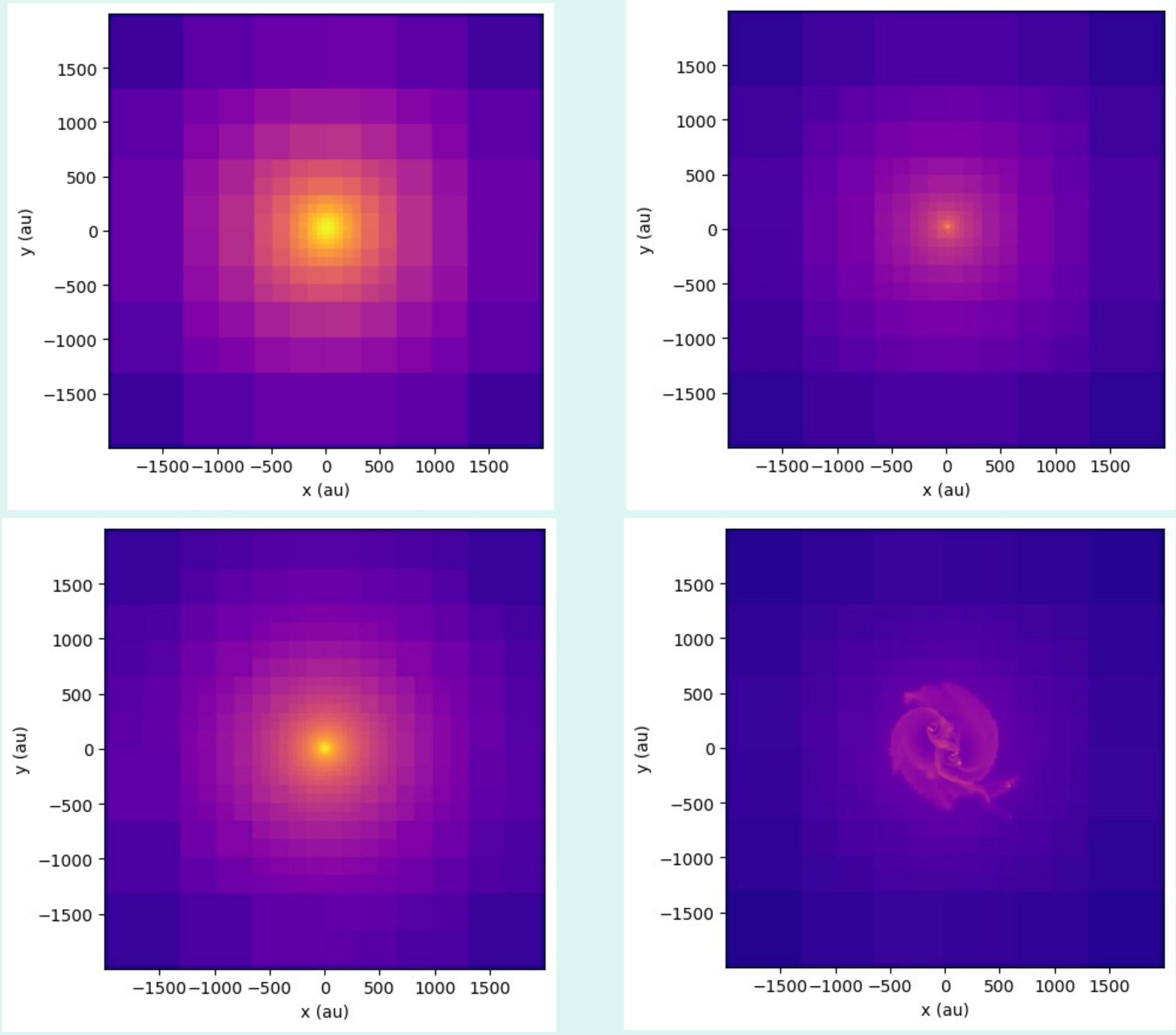
	Extension of the aBplfit technique (Veltchev et al. 2019)
	• Averaged PDFs (over varied total number of bins) • *Input parameters:* lower cutoff, upper cutoff, range of variation of the total number of bins. • *Output (PLT) parameters:* slope, deviation point

Development of multiple PLTs

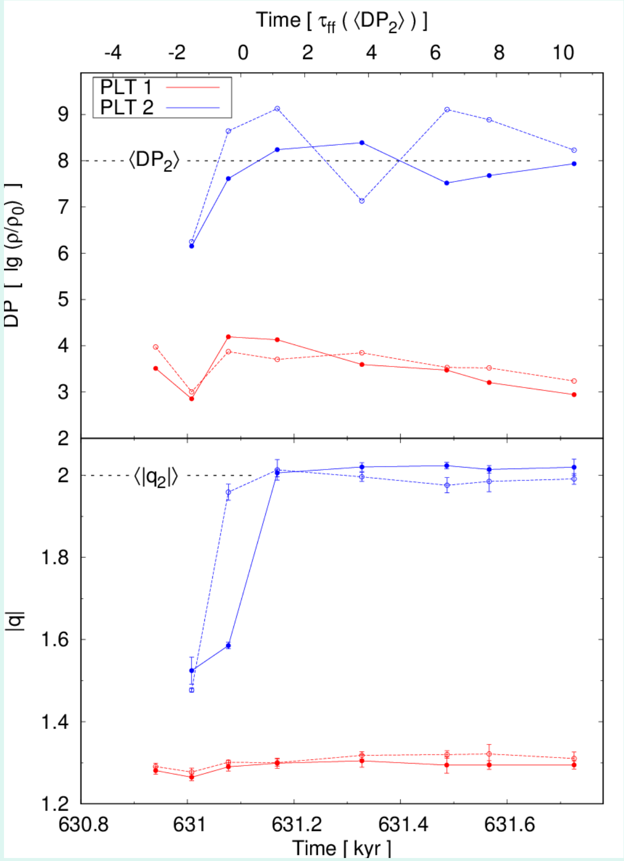


$t = t_{SF}$ $t = t_{end}$



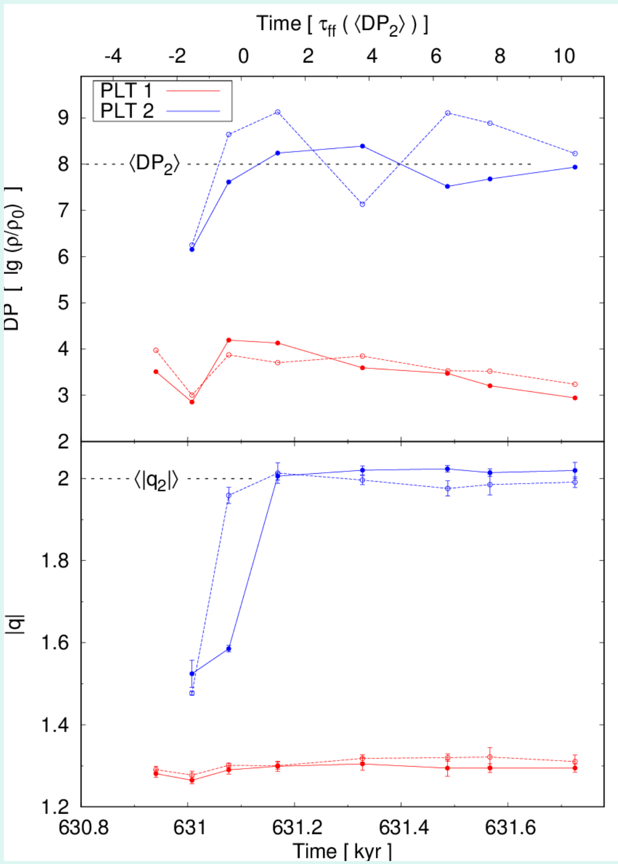
Evolution of the PLTs in *ρ*-PDFs

PI

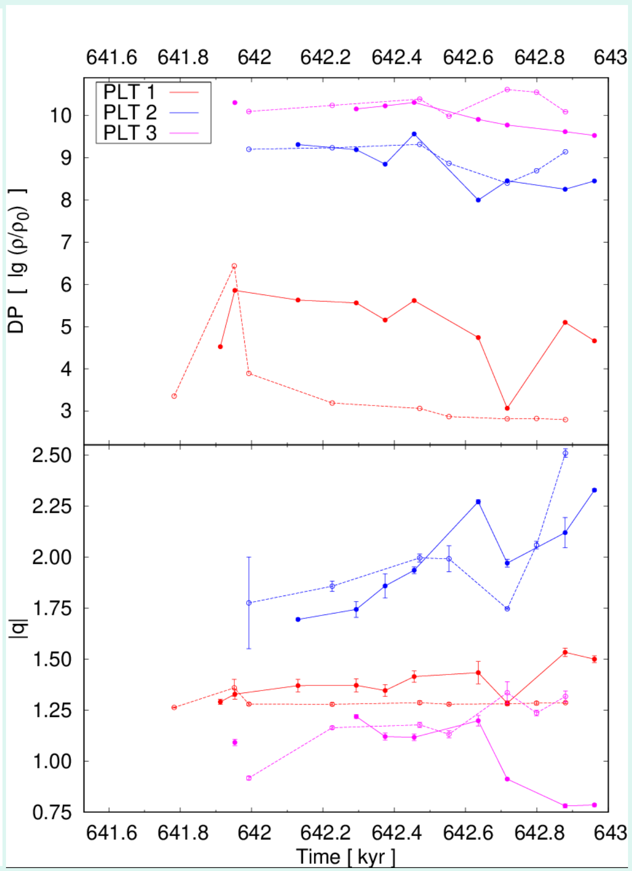


Evolution of the PLTs in ρ-PDFs

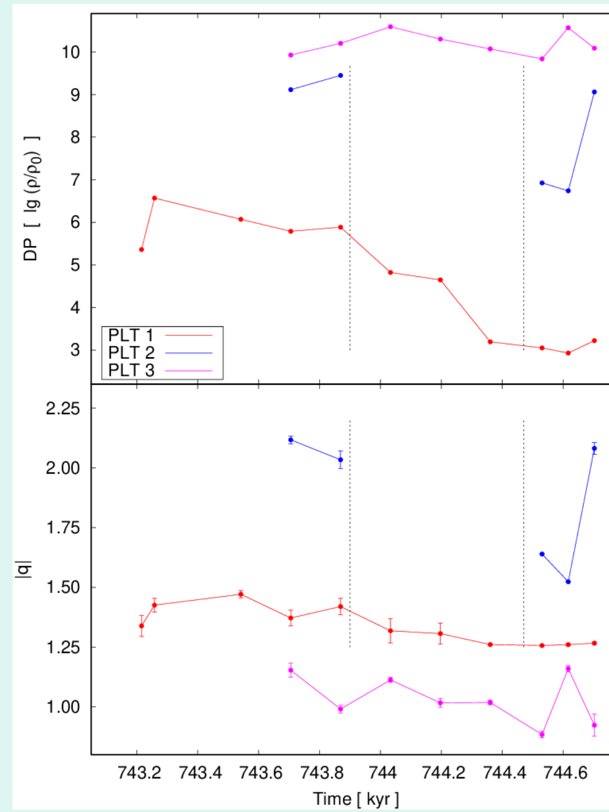
PI



RO, β=0.01



RO, β=0.10

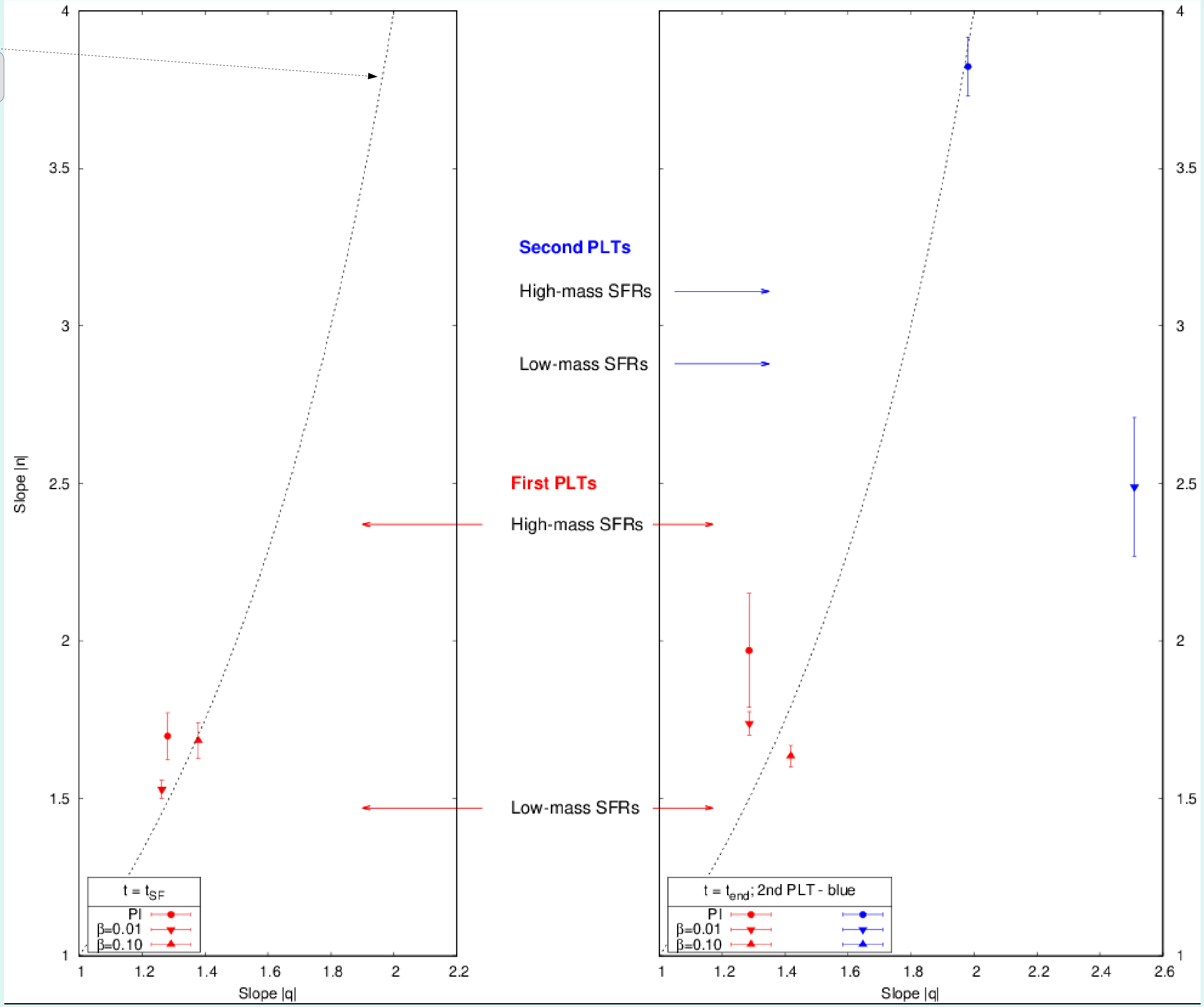


Evolution of the PLTs in *N*-PDFs

$$n = 2q/(3+q)$$

Donkov, Veltchev & Klessen (2017)

- Relation between the exponents of the PLTs in *ρ*– and *N*-PDF.
- Based on the assumption for spherical symmetry and density profile of power-law type.



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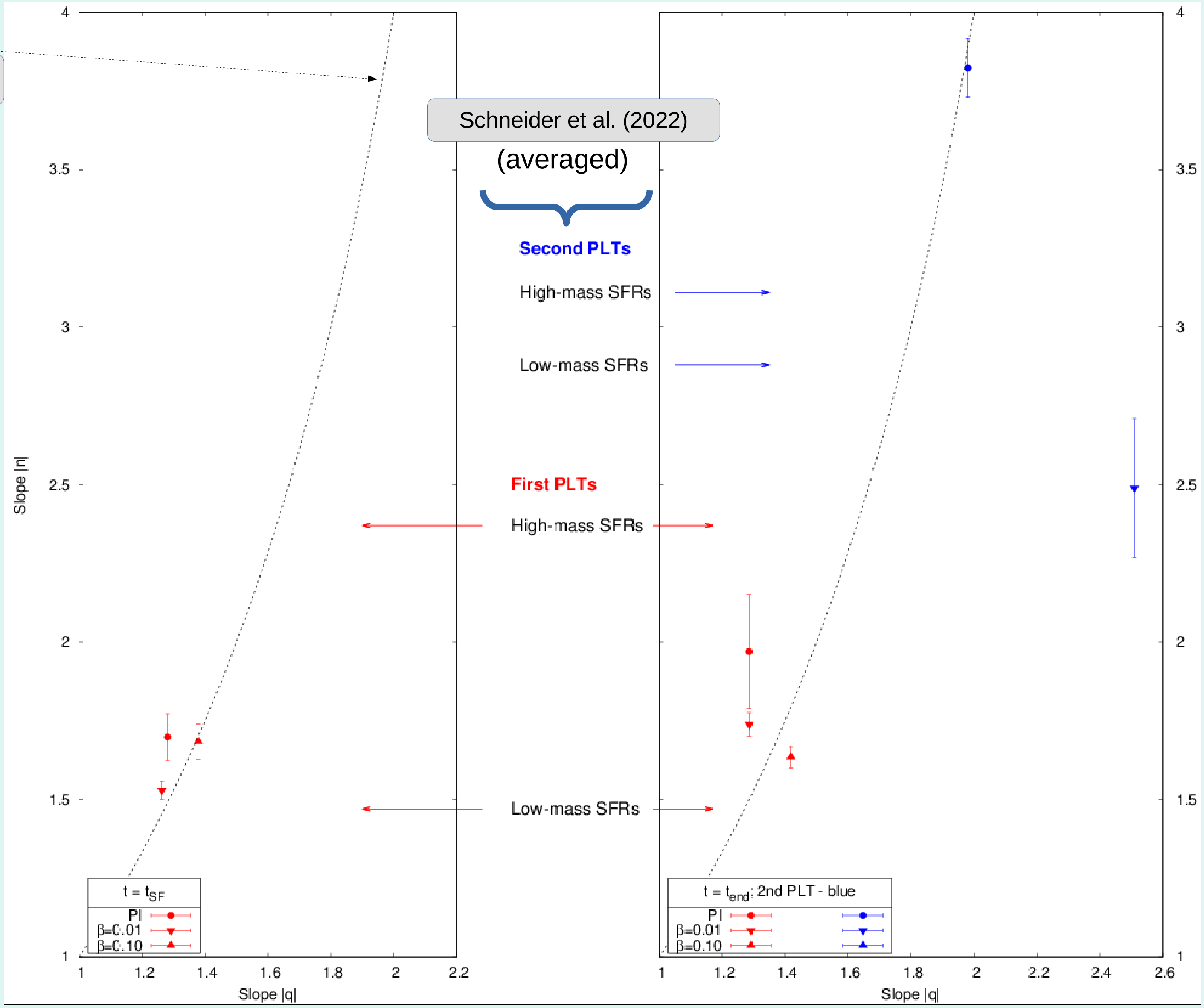
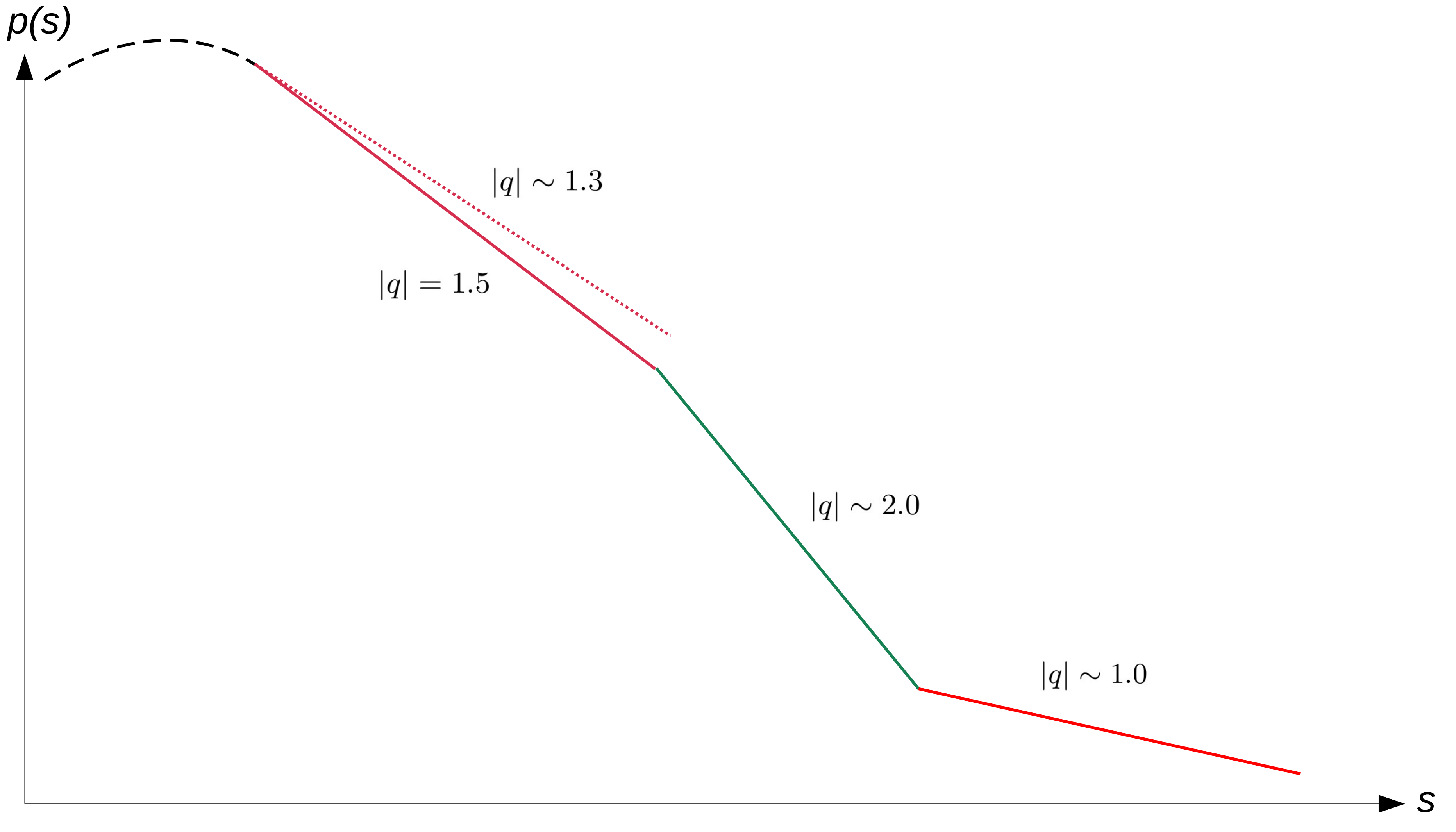
- Relation between the exponents of the PLTs in *ρ*– and *N*-PDF.
- Based on the assumption for spherical symmetry and density profile of power-law type.



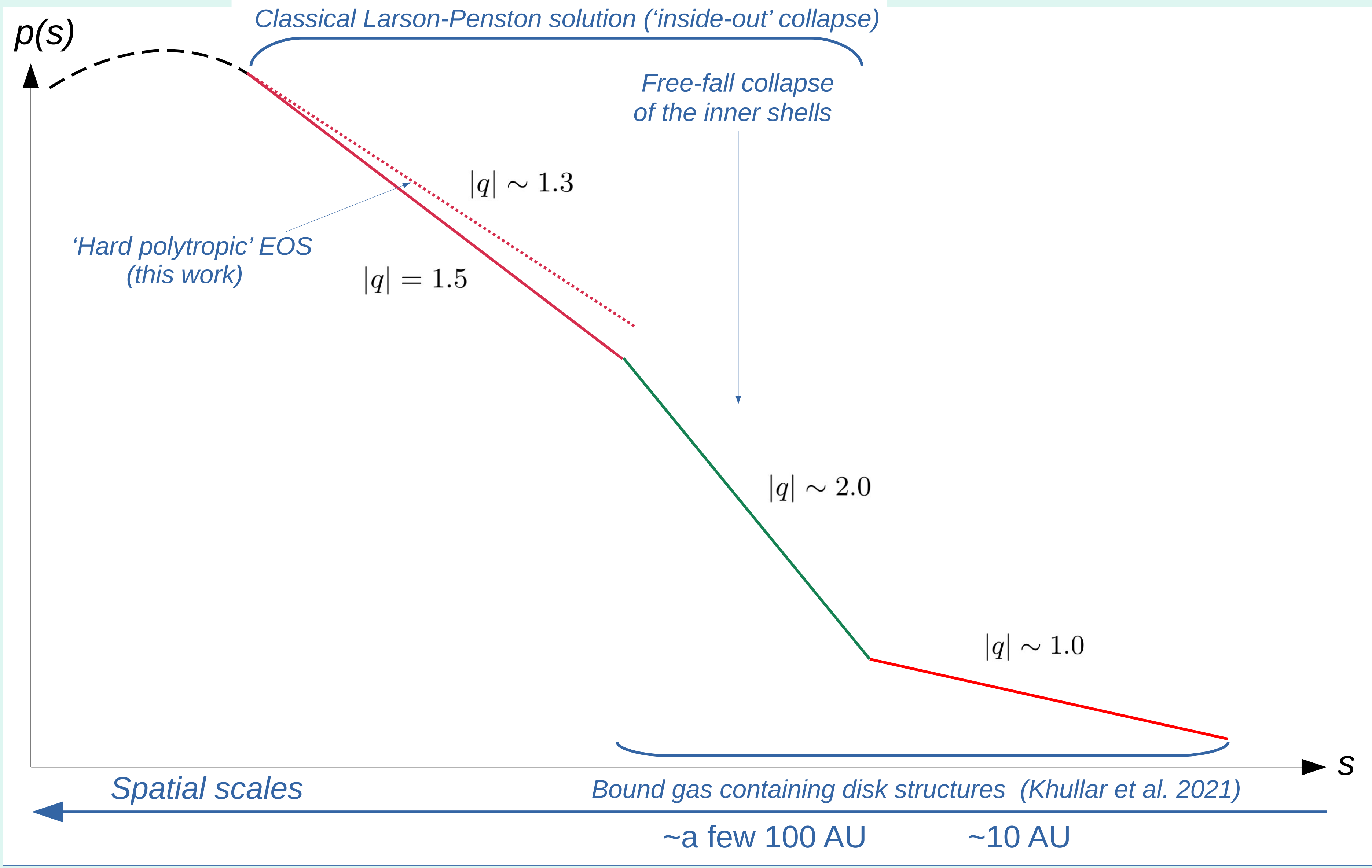
Results: evolution of PLTs in SF clouds

- Similarities between all runs (pure infall and with rotational support)
  - Emergence of PLT 1 at $t \sim t_{SF}$; it retains its slope ($q_1 \sim -1.3$) within many free-fall times
  - Emergence (shortly after $t \sim t_{SF}$) and development of PLT 2 at the high-density end of the PDF, with a typical value $q_2 \sim -2.0$
  - *(For the runs with rotational support)* Emergence of PLT 3 at the very high-density end (i.e. very small spatial scales) whose slope varies around a typical value $q_3 \sim -1.0$
  - Relation between the PLT 1 slopes in ρ- and N-PDF corresponds to a spherically symmetric model with a radial PL density profile. They are in general agreement with the recent observations of regions of various SF activity (Schneider et al. 2022).
- Differences:
  - No PLT 3 in the pure-infall runs
  - Unstable PLT 2 in the RO runs; it disappears occasionally for β=0.10

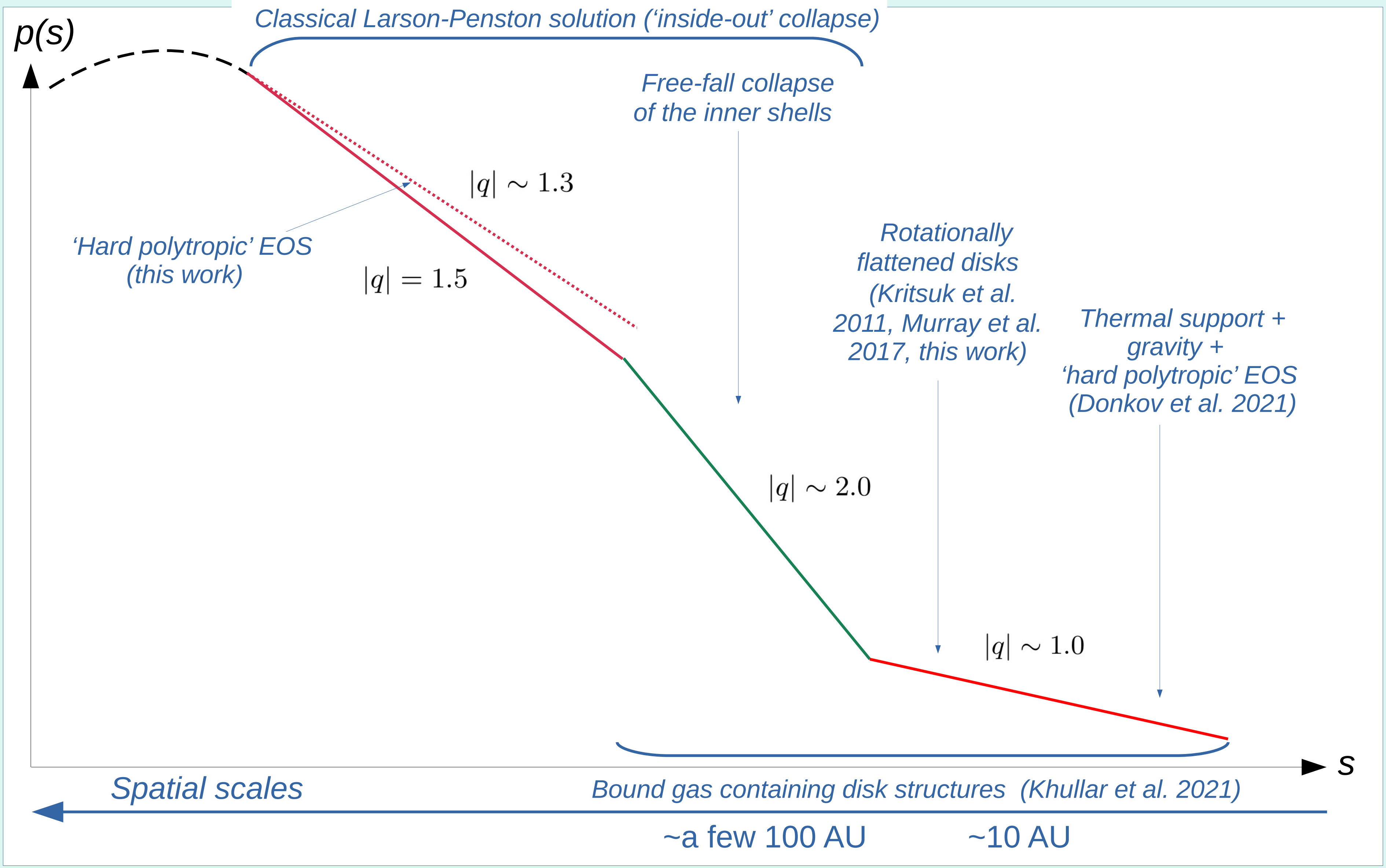
Contribution: multiple PLTs in SF clouds/clumps



Contribution: multiple PLTs in SF clouds/clumps



Contribution: multiple PLTs in SF clouds/clumps



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